Infrared astronomy from Antarctica Mapping of the aliphatic component of the Galaxy

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Hong Kong March 2015

Absorption/scattering dependence on color

Solids in the universe

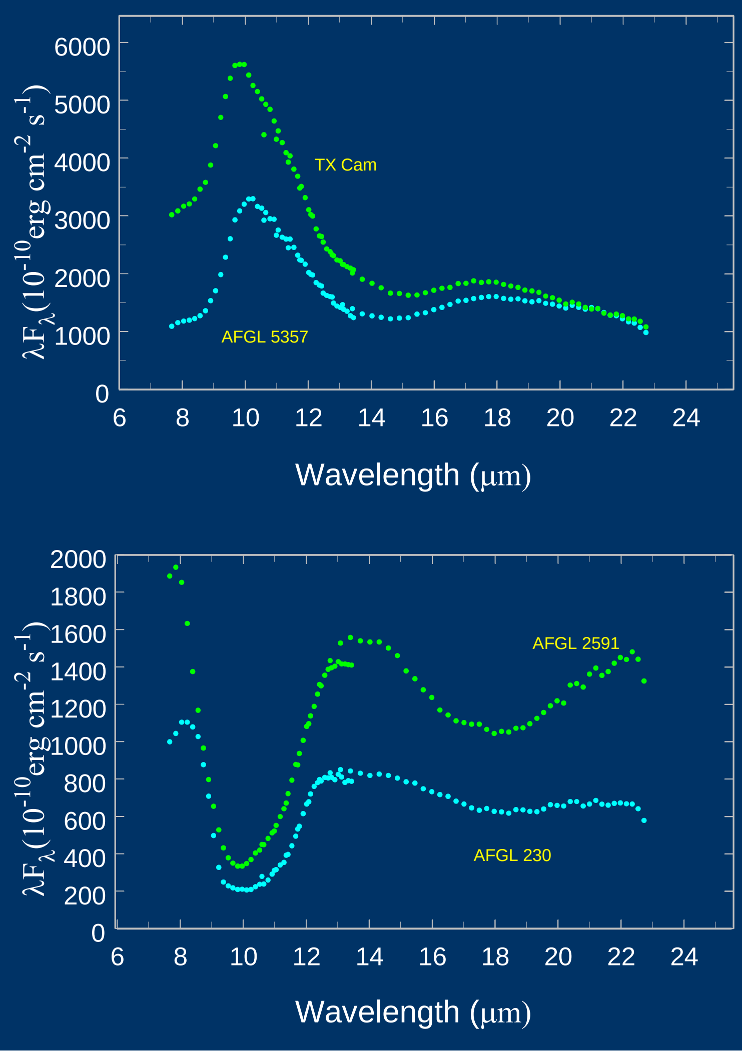
- Early models based on the shape of the extinction curve generally assumed graphite, iron, ice, etc.
- Crystalline solids have highly ordered lattice structure and have optically active lattice vibrational modes
- Amorphous solids have many optically active modes
- Large organic solids have many electronic states and their absorption bands are very broad. Macromolecular organic solids can be identified by their visible colors, albedo, and *infrared bands*.
- Carbonaceous solids: ices, graphite, diamonds, fullerenes, amorphous carbons, hydrogenated amorphous carbons, ….

Condensation of solids:
Amorphous Silicates

- IRAS Low Resolution Spectrometer
- Over 11,000 spectra processed and classified (Kwok et al. 1997)
- Seen in both emission and self absorption

9.7 µm: Si-O stretch
18 µm: O-Si-O bend

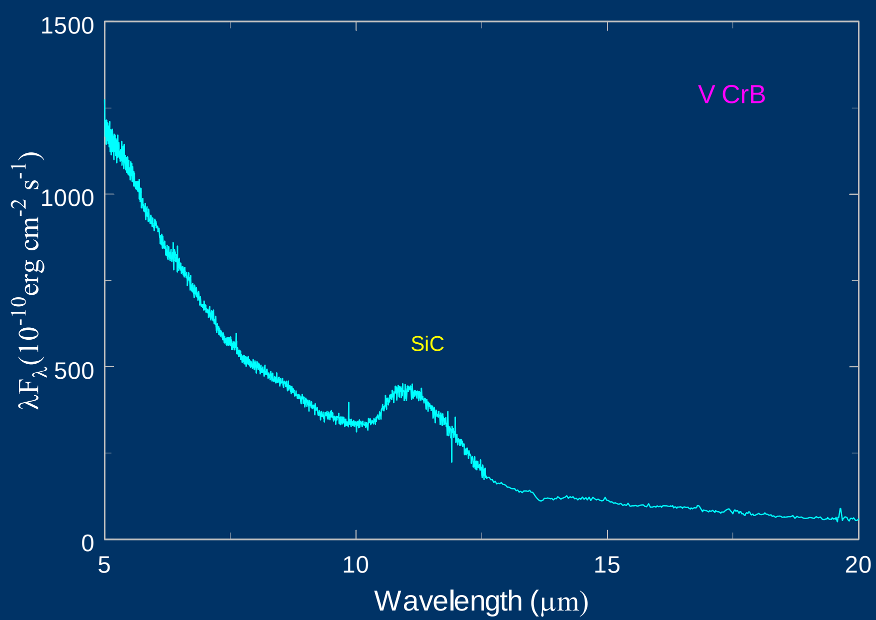
IRAS LRS



Direct condensation from gas to solid results in disordered structure

Silicon Carbide

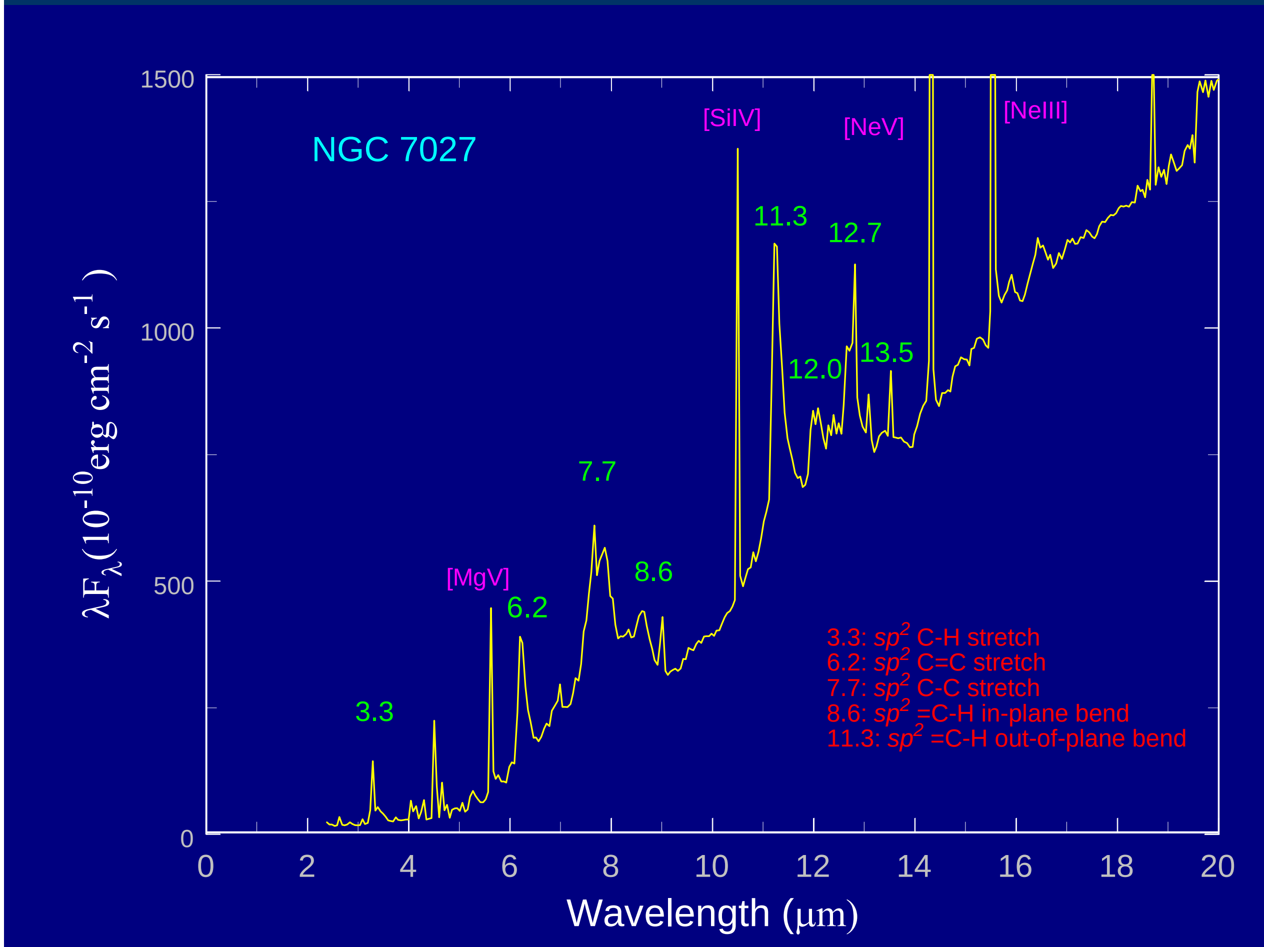
- 4000 stars detected to have amorphous silicates by IRAS LRS
- 700 stars detected in SiC
- M stars: silicates, C stars: SiC



Do dust in planetary nebulae has the same chemical composition as dust in AGB stars?

Unidentified infrared emission bands

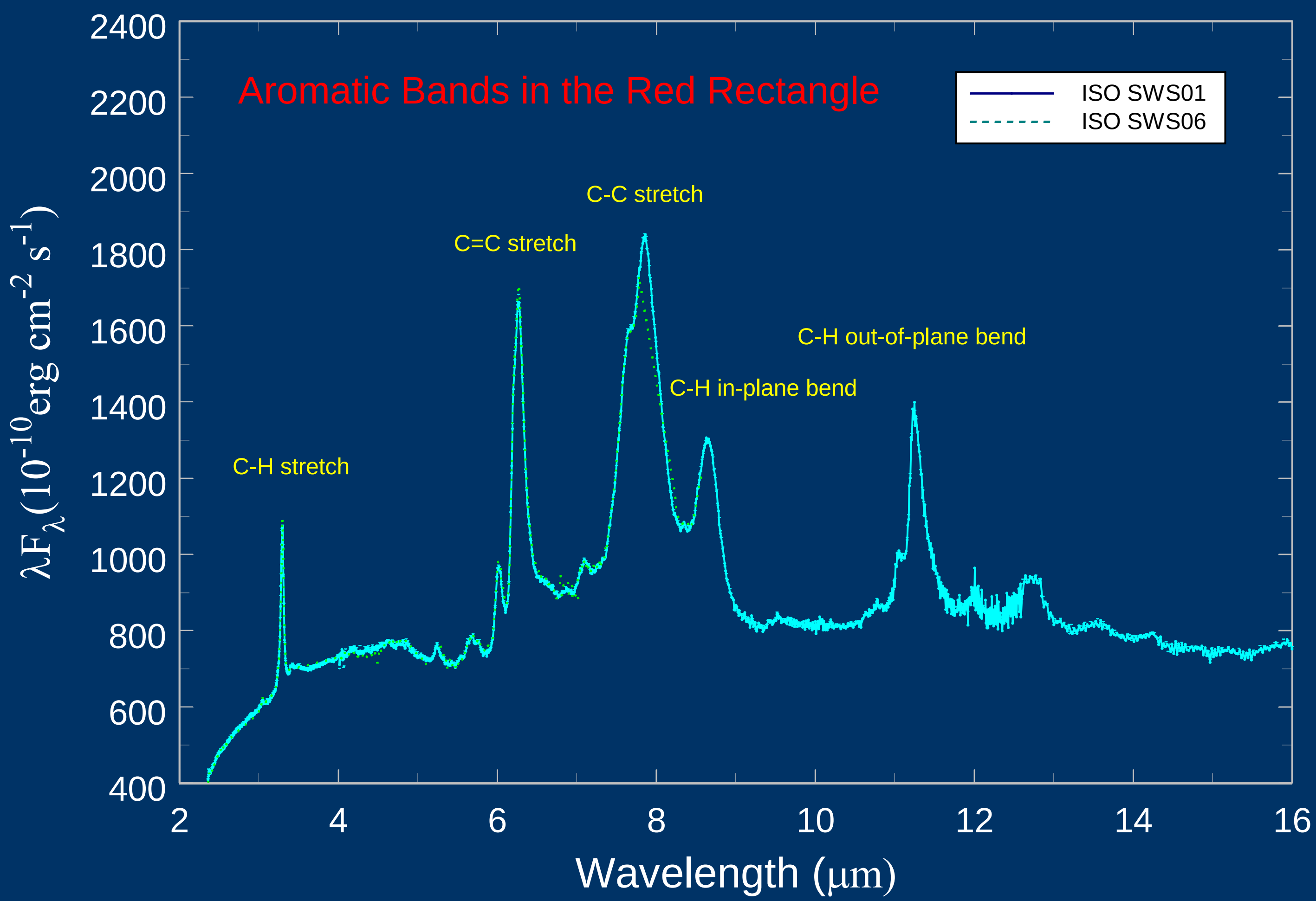
(Russell et al. 1977)

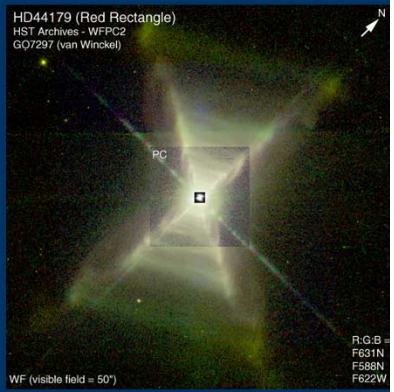


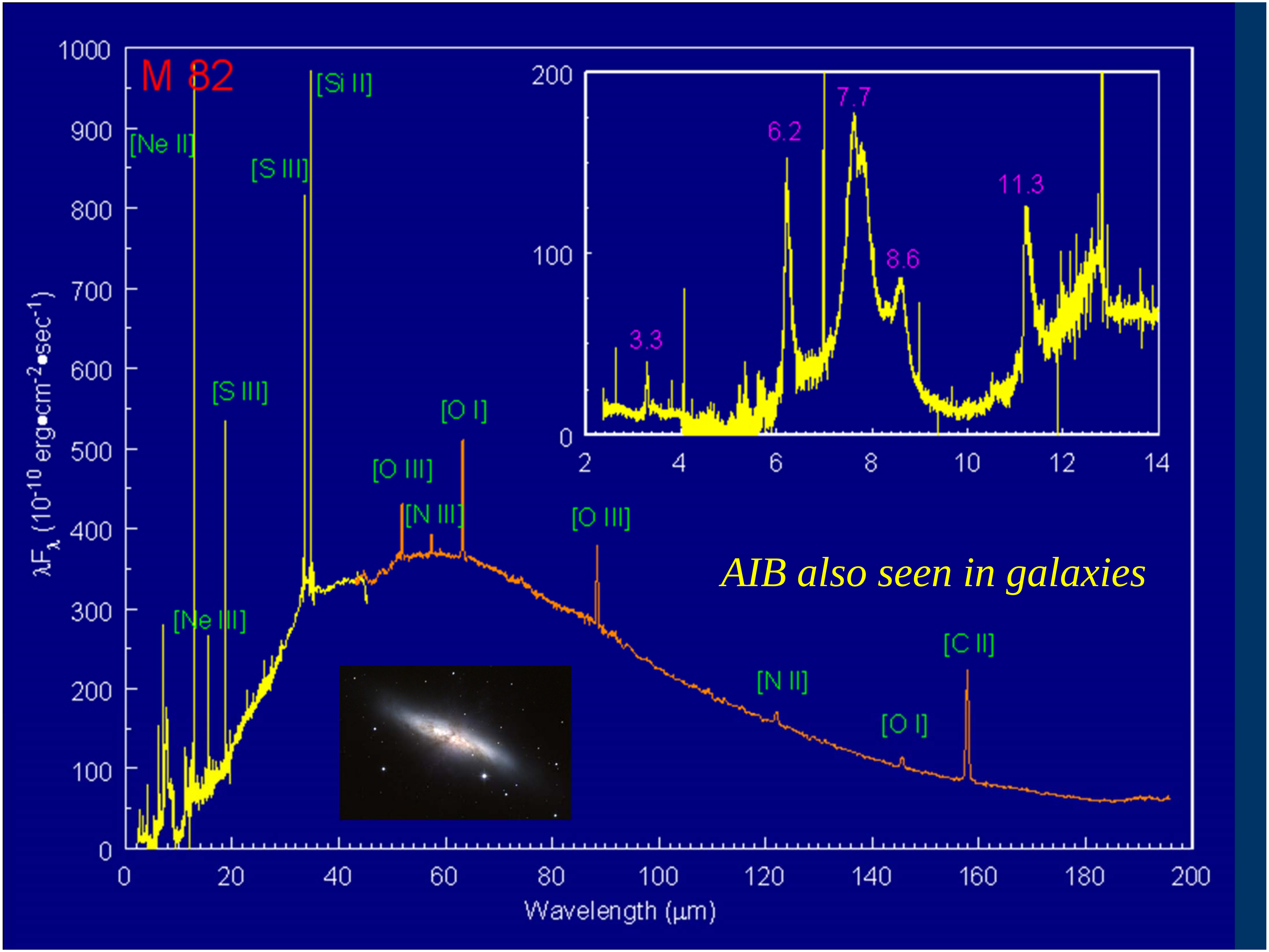
Aromatic nature first proposed by: Knacke 1977, Duley & Williams 1979, 1981; Puetter et al. 1979

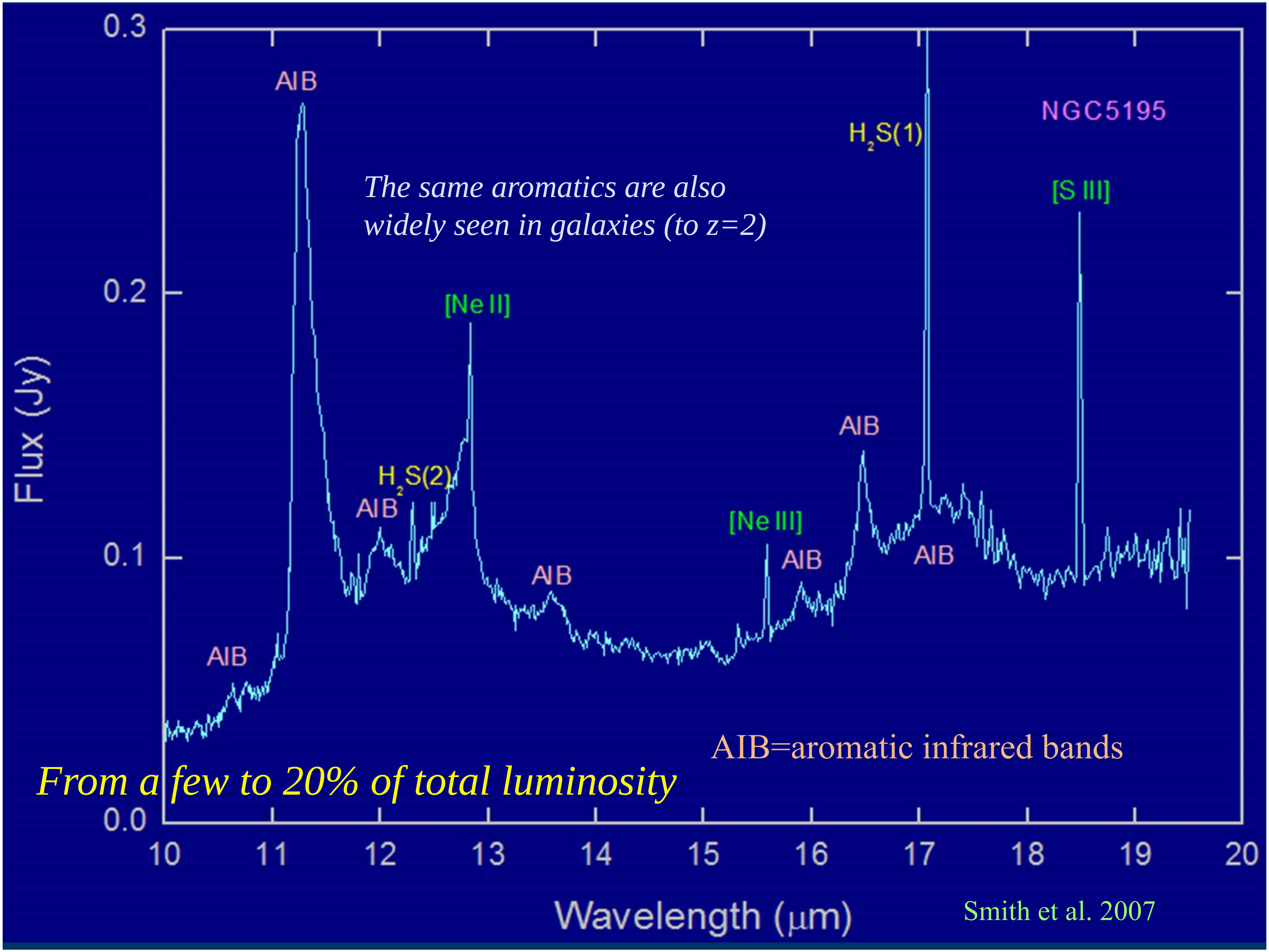
Stretching and bending modes of aromatic compounds

AIB seen in reflection nebulae, HII regions, diffuse ISM, and galaxies



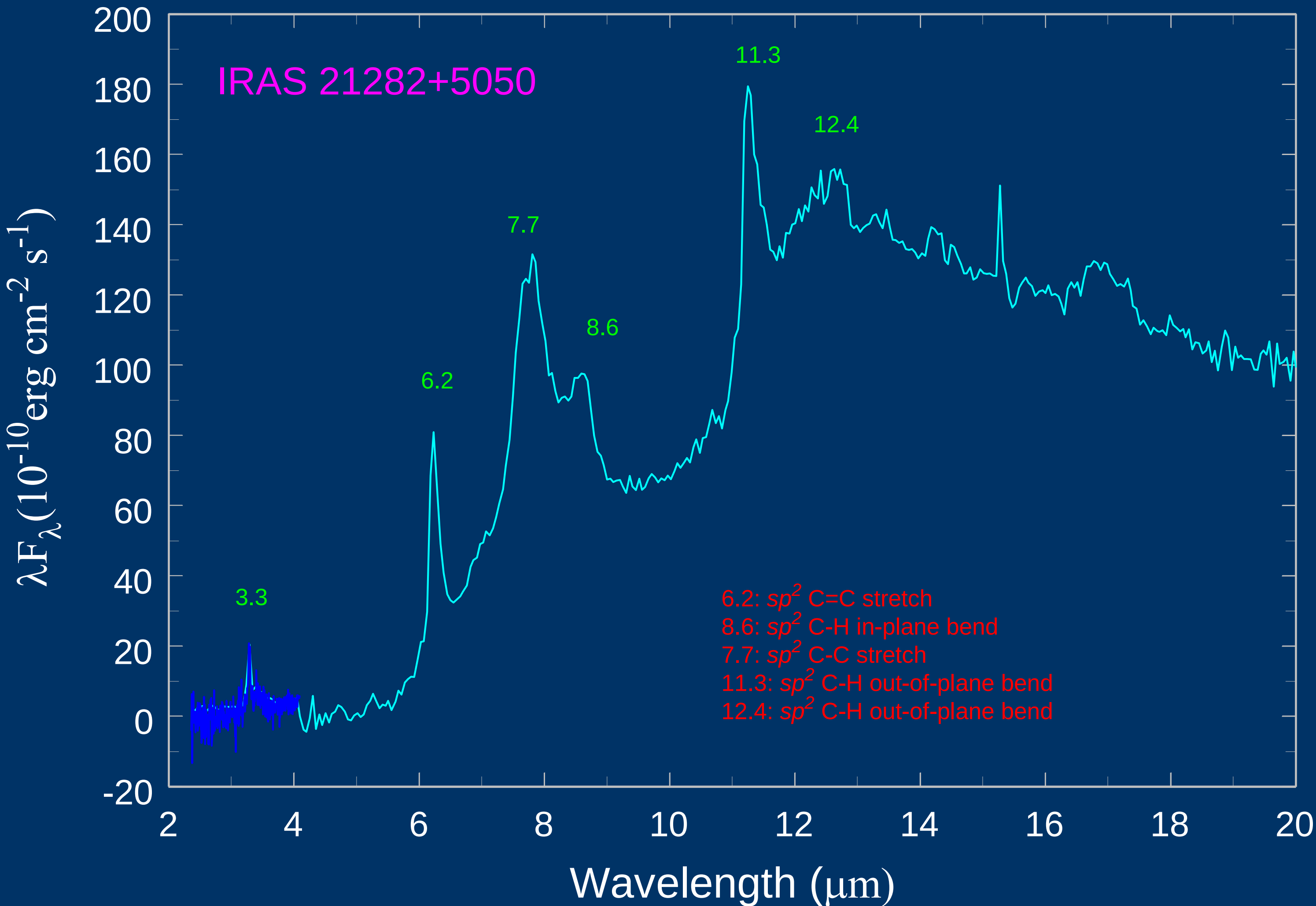



M 82
[Ne II]
[S III]
[Si II]
[S III]
[Ne III]
[O III]
[N III]
[O I]
[O III]
[N II]
[O I]
[C II]
3.3
6.2
7.7
8.6
11.3
λF_λ (10^-10 erg•cm^-2•sec^-1)
Wavelength (μm)
AIB also seen in galaxies



The same aromatics are also widely seen in galaxies (to z=2)
From a few to 20% of total luminosity
AIB=aromatic infrared bands
NGC5195
AIB
[Ne II]
H2S(2)
[Ne III]
H2S(1)
[S III]
Flux (Jy)
0.0
0.1
0.2
0.3
10
11
12
13
14
15
16
17
18
19
20
Wavelength (μm)
Smith et al. 2007

AIB are detected in many planetary nebulae. Since the carrier is synthesized in situ, PN are the best objects to study their origins



A young PN

When are the aromatic compounds synthesized?

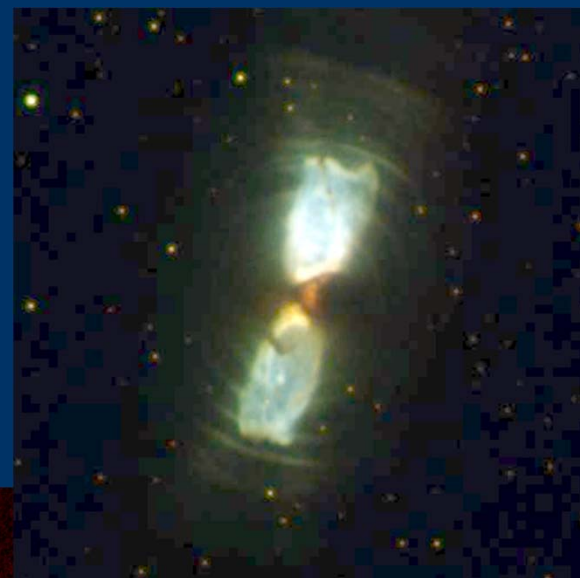
- AIB features not seen in the progenitor AGB stars (dynamical age ~10^4 yr)
- AIB features are strong in young planetary nebulae (age <10^4 yr)
- Have to study the missing link between AGB and PN phases

Proto-planetary nebulae

- Objects in transition between AGB and PN stages *(about several thousand years)*
- Difficult to identify because the nebulae of PPN are neutral and do not have emission lines
- ~30 PPN were discovered from their IR properties as the result of follow up of IRAS survey (Kwok 1993, *Ann. Rev. Astr. Ap.*, 31, 63)

PPN as imaged by the HST

The Silkworm Nebula



The Cotton Candy Nebula

The Walnut Nebula

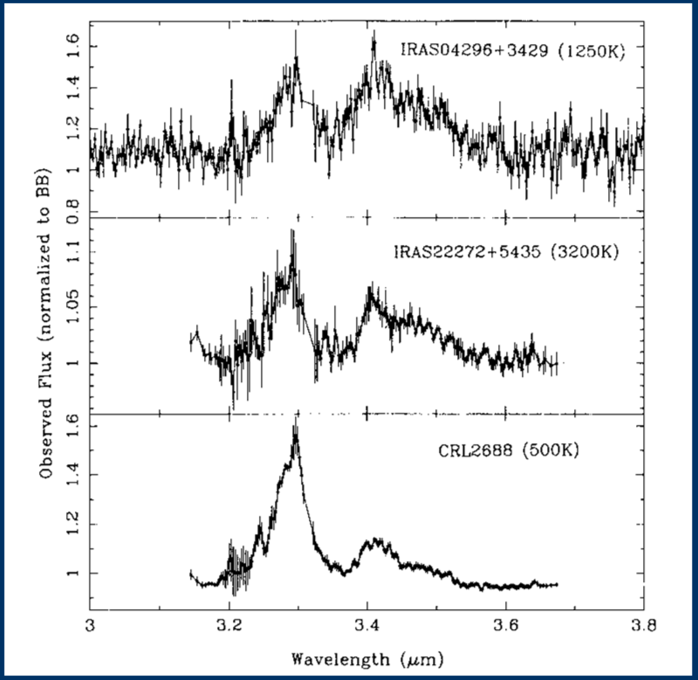
The Water Lily Nebula

The Spindle Nebula

Reflected starlight, not emission!

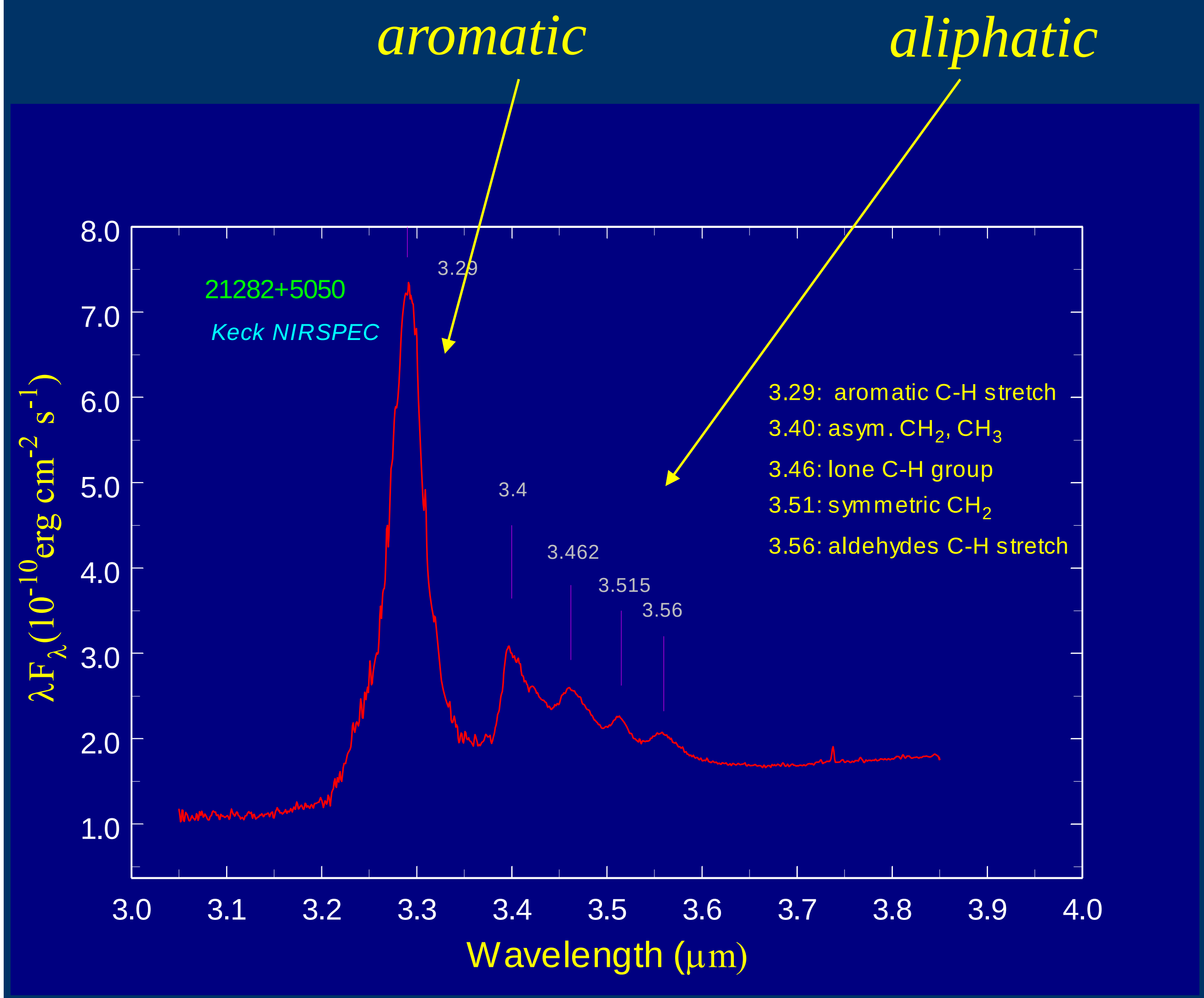
3.4 μm aliphatic C-H stretch

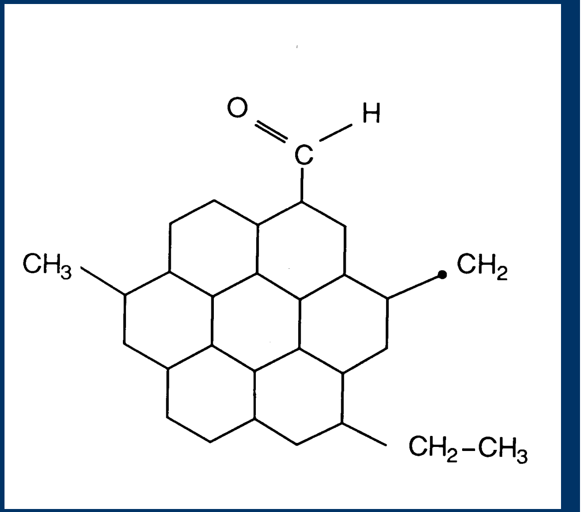
- 3.38 μm: asymmetric CH_3
- 3.42 μm: asymmetric CH_2
- 3.46 μm: lone C-H group
- 3.49 μm: symmetric CH_3
- 3.51 μm: asymmetric CH_2

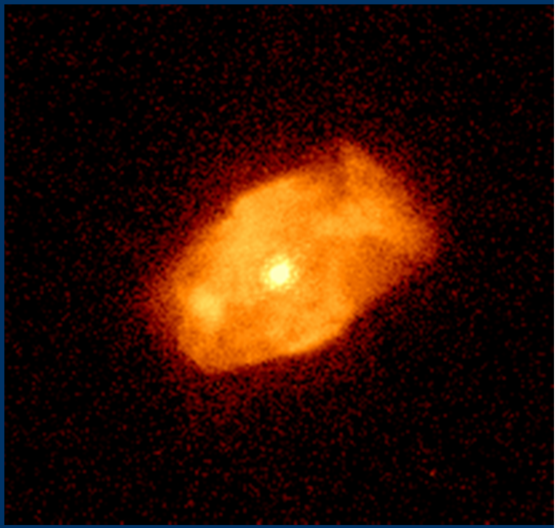


Infrared spectroscopy reveals aliphatic features

Aliphatic sidegroups



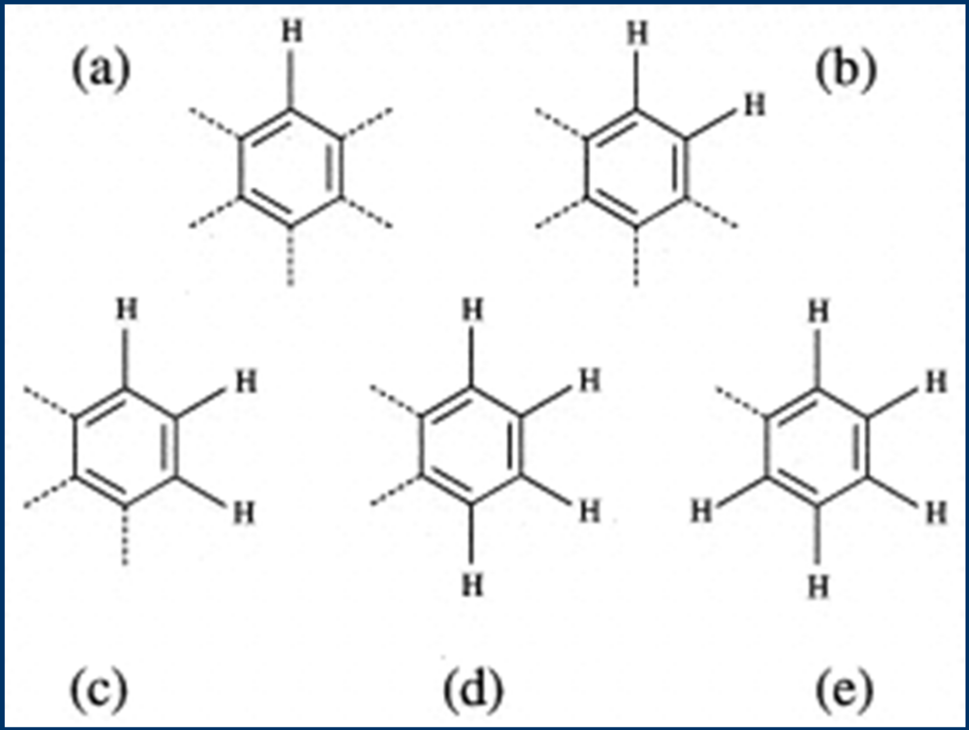




Number of CH groups in aromatic rings

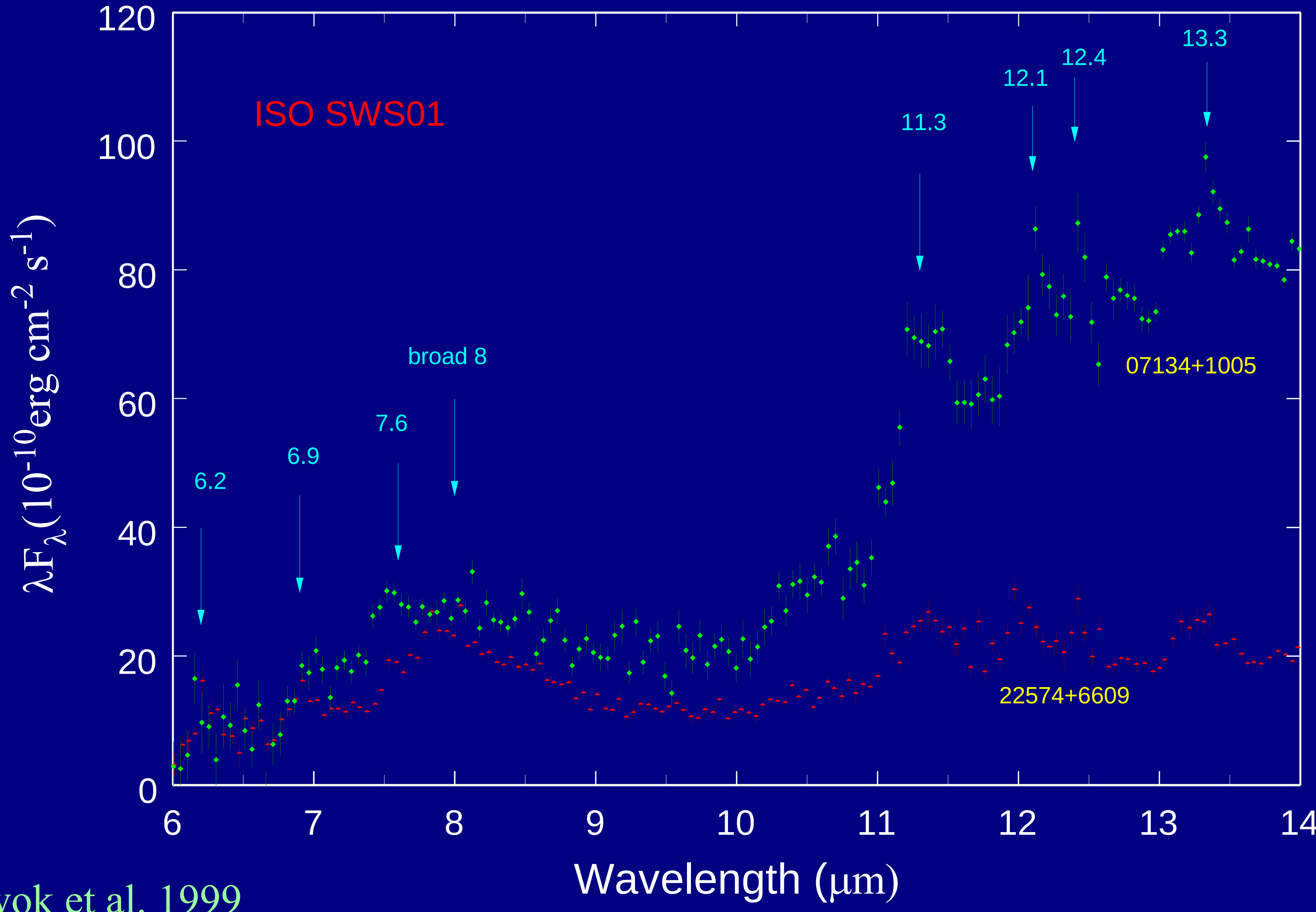
Their out of plane bending mode frequencies are different

- Solo: 11.1-11.6 μm
- Duo: 11.6-12.5 μm
- Trio:12.4-13.3 μm
- Quarto: 13-13.6 μm



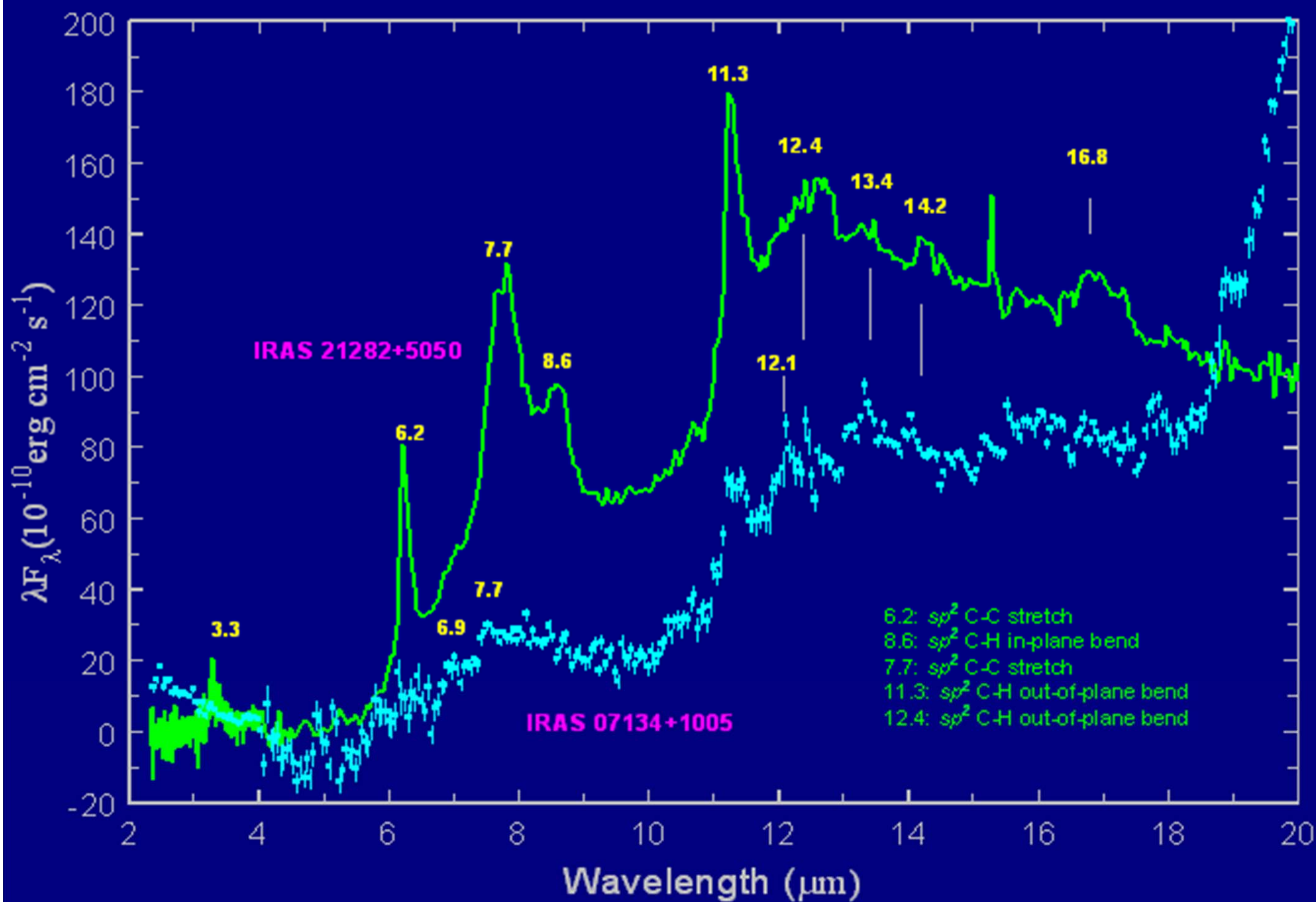
Hugdins and Allamandola 1999

The number of rings in each aromatic unit is small

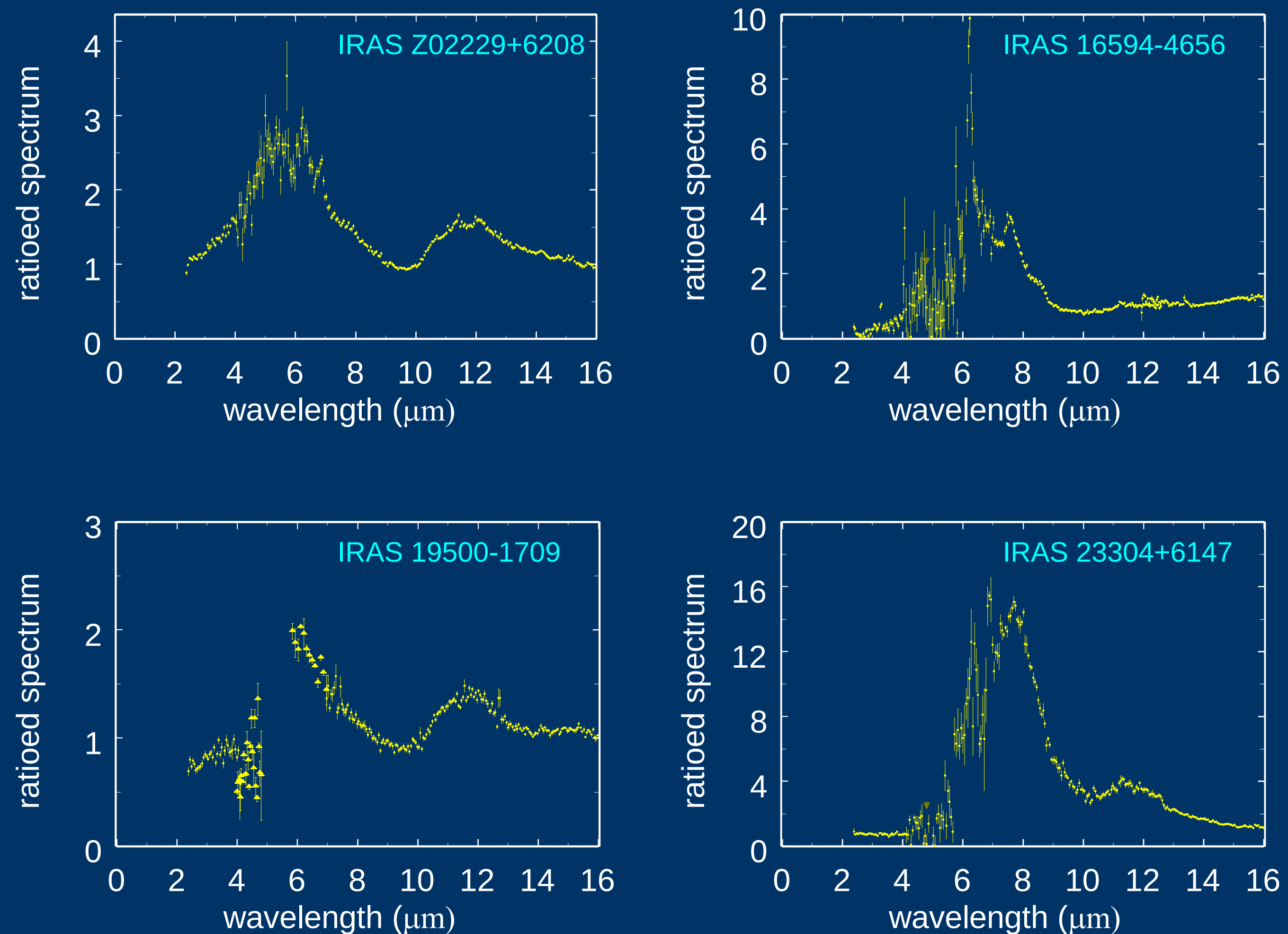


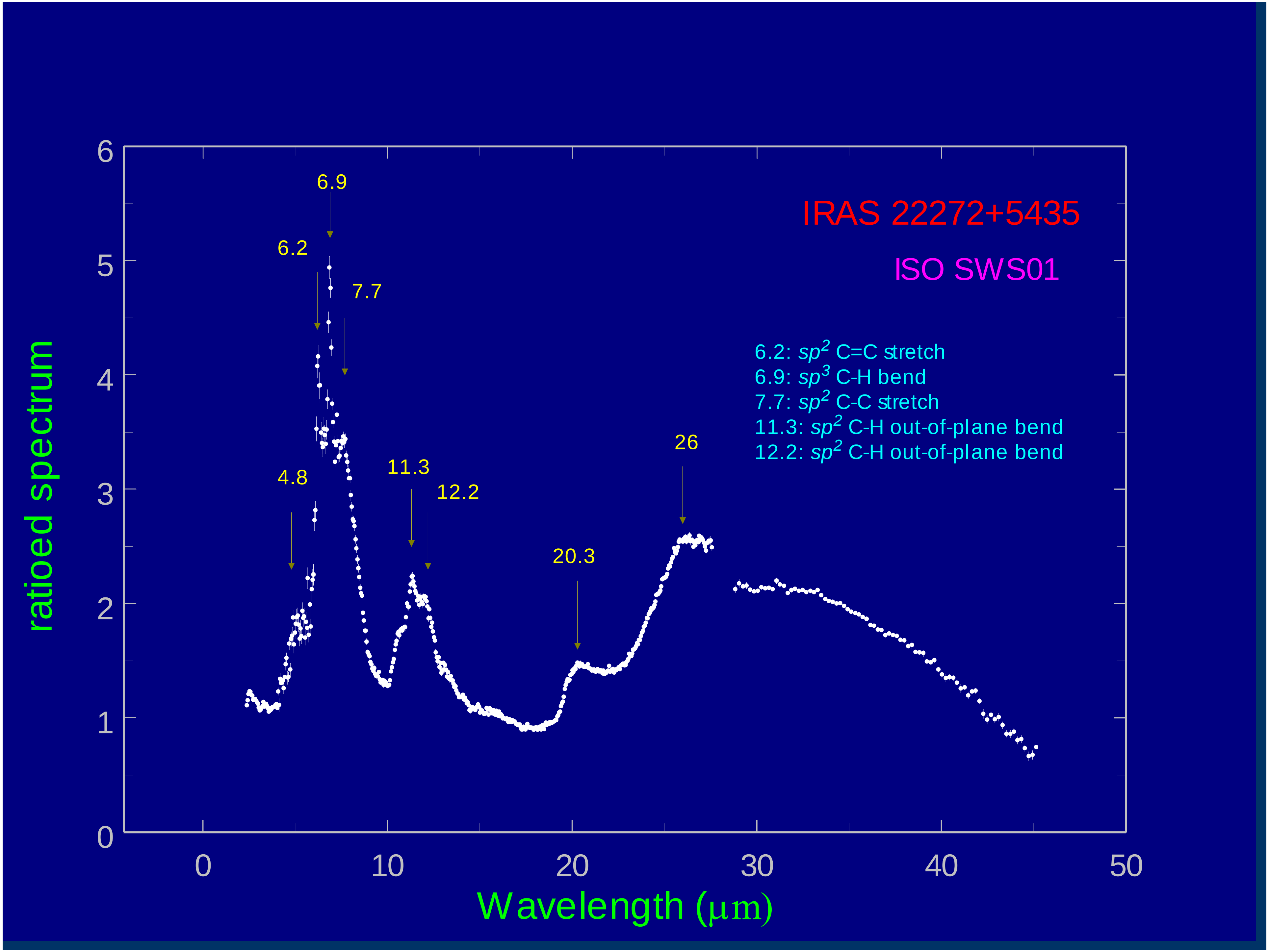
Kwok et al. 1999

Aromatic and aliphatic features in PPN



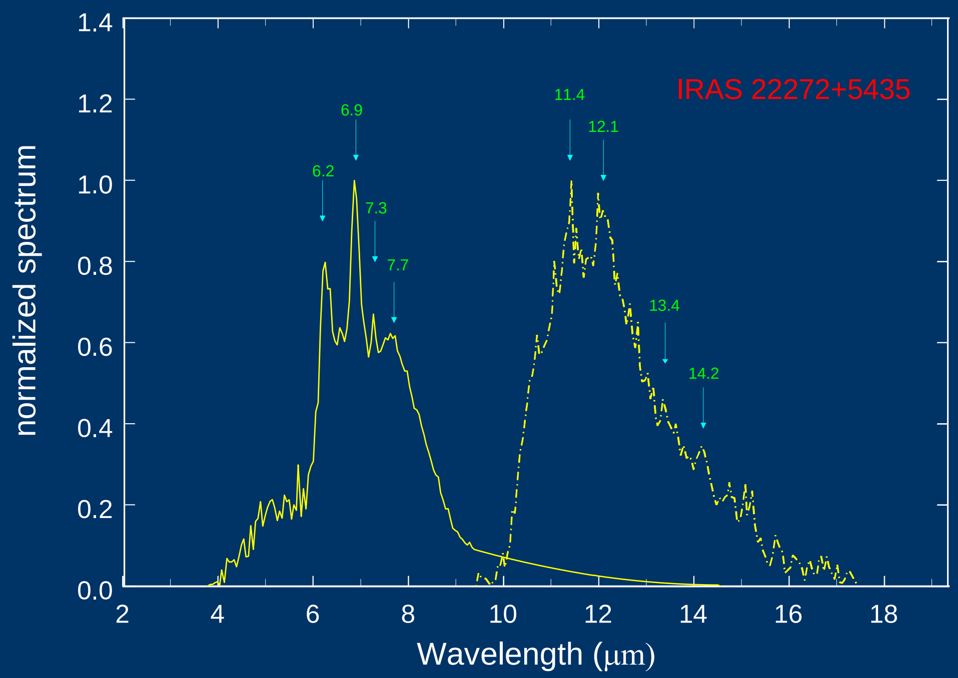
Broad emission plateaus

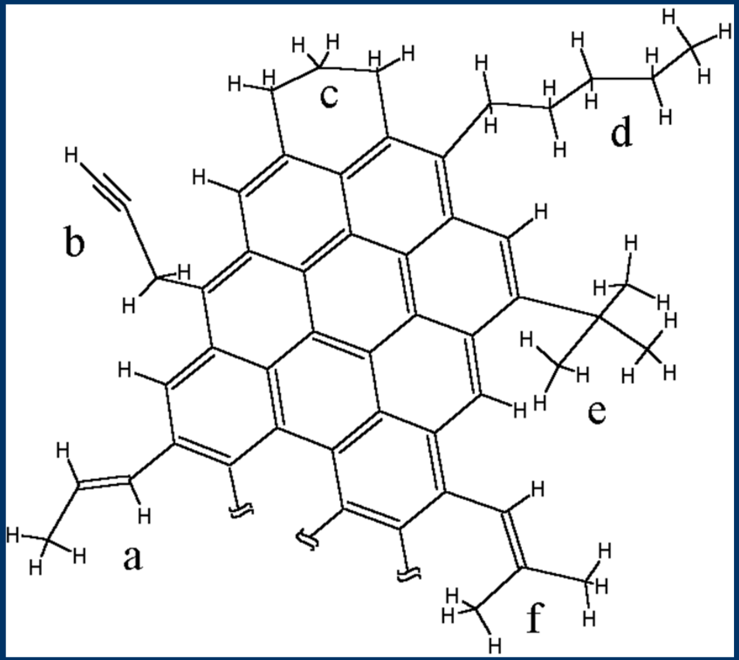




IRAS 22272+5435
ISO SWS01
6.2: sp2 C=C stretch
6.9: sp3 C-H bend
7.7: sp2 C-C stretch
11.3: sp2 C-H out-of-plane bend
12.2: sp2 C-H out-of-plane bend
6.9
6.2
7.7
4.8
11.3
12.2
20.3
26
ratioed spectrum
Wavelength (μm)
0
1
2
3
4
5
6
0
10
20
30
40
50

Aliphatic in-plane and out-of-plane bending modes





Kwok et al. 2001

- 8μm plateau: $-CH_3$ (7.25 μm), $-C(CH_3)_3$ (8.16 μm, "e"), $=(CH_3)_2$ (8.6 μm, "f")
- 12 μm plateau: C-H out-of-plane bending modes of alkene ("a", "b"), cyclic alkanes (9.5-11.5 μm, "c"), long chains of $-CH_2-$ groups (13.9 μm, "d").

Advantages of circumstellar chemistry

- Single energy source
- Simple geometry
- Well-determined physical environment (density $\rho(r)$, temperature $T(r)$, radiation background $I(r)$)
- Chemical time scale defined by dynamical time scale (AGB: 10^4 yr, PPN:10^3 yr, PN: 10^4 yr)

The Unidentified Infrared Emission bands phenomenon

- Aromatic features: 3.3, 6.2, 7.7, 8.6, and 11.3 μm
- Aliphatic features: 3.4 and 6.9 μm
- Features at 15.8, 16.4, 17.4, 17.8, and 18.9 μm (in PPN, Kwok et al. 1999, in reflection nebulae, Sellgren et al. 2007, in galaxies, Sturm et al. 2000)
- Broad plateau features at 8, 12, and 17 μm.

What is the chemical structure of the carrier?

The PAH hypothesis

(Allamandola et al. 1989, Puget & Léger 1989)

- the UIE features are the result of infrared fluorescence from small (~50 C atoms) gas-phase PAH molecules being pumped by far-ultraviolet photons (Tielens 2008, *Ann. Rev. Astr. Ap.*, **46**, 289)
- The central argument for the PAH hypothesis is that single-photon excitation of PAH molecules can account for the 12 μm excess emission observed in cirrus clouds in the diffuse interstellar medium by *IRAS* (Sellgren 1984, 2001).

Problems with the PAH model

- PAH molecules have well-defined sharp features but the UIE features are broad
- PAHs primarily excited by UV, with little absorption in the visible, but UIE features are seen in PPN and reflection nebulae with no UV radiation
- The strong and narrow predicted gas phase features in the UV are not seen in interstellar extinction curves → upper limits of 10^{-10}-10^{-8} (Clayton et al. 2003, Salama et al. 2011, Gredel et al. 2011)
- No specific PAH molecules have been detected in spite of the fact that the vibrational and rotational frequencies are well known

Expected from IR: $3\text{x}10^{-7}$ (Tielens 2008)

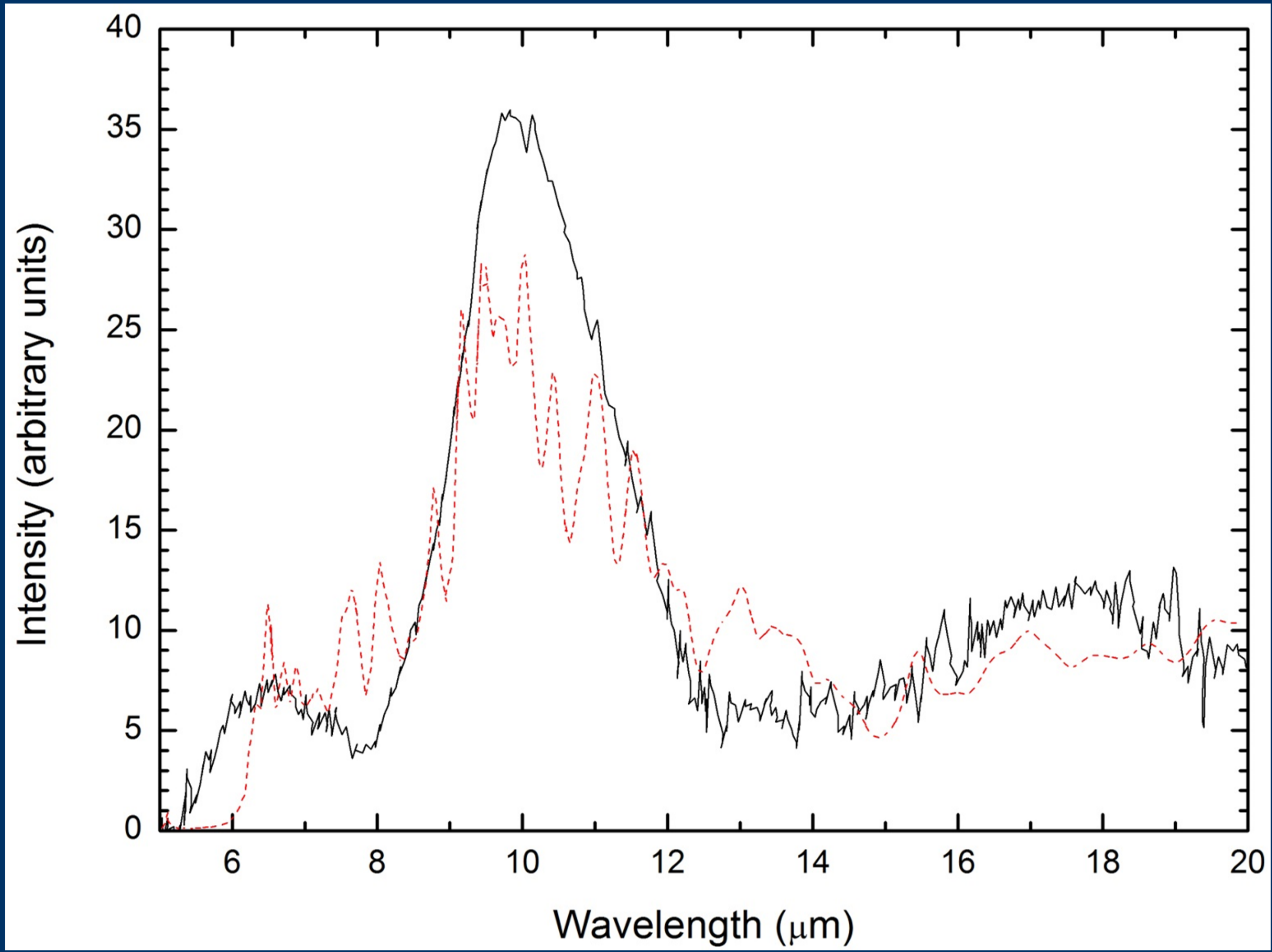
Problems with the PAH model

- "No PAH emission spectrum has been able to reproduce the UIE spectrum w.r.t. either band positions or relative intensities" (Schlemmer et al. 1994, Cook et al. 1996, Cook & Saykally 1998, Wagner et al. 2000)
- The shapes and peak wavelengths of UIE features are independent of temperature of exciting star
- In order to fit the astronomical observations, the PAH model has to appeal to a mixture of PAH of different sizes, structures (compact, linear, branched) and ionization states, as well as artificial broad intrinsic line profiles (Cami 2011).

Fitting of UIE by PAHs

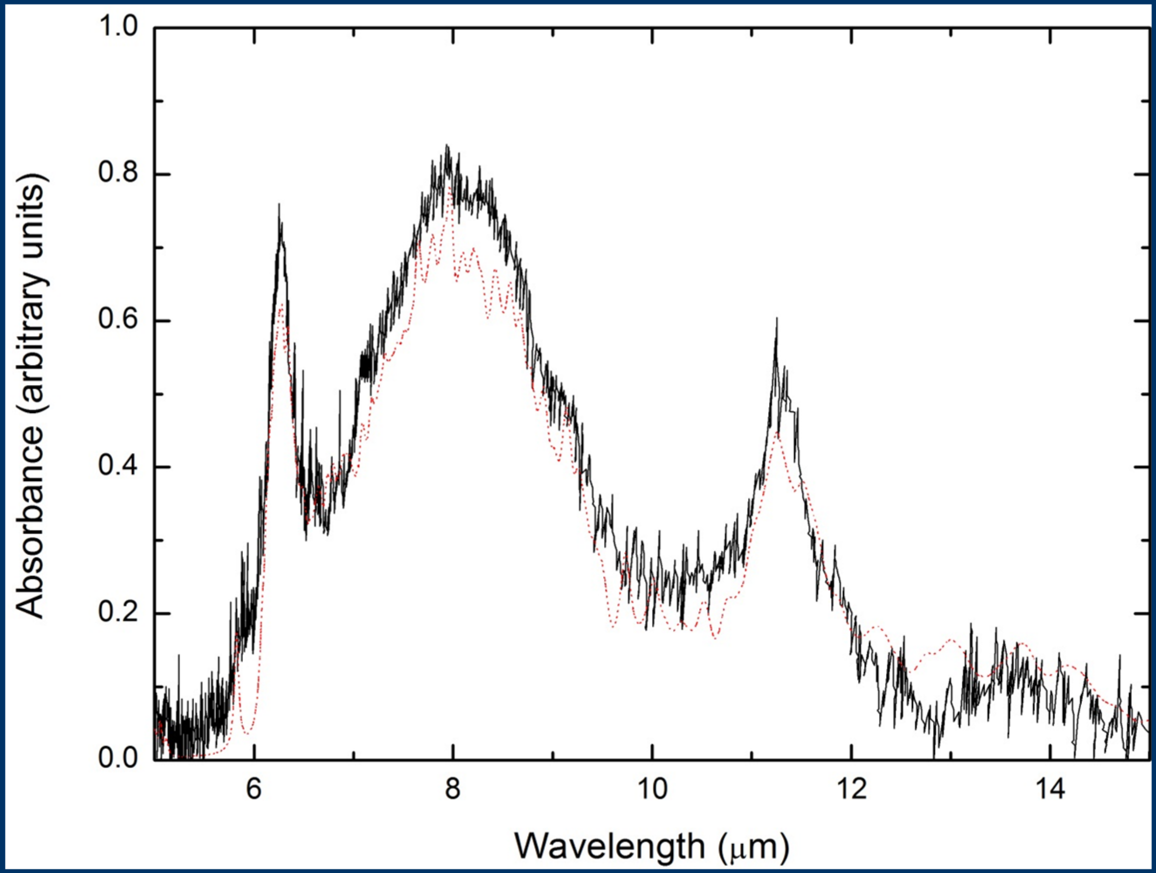
- NASA Ames PAH database and fitting routines (Boersma et al. 2014).
- 700 computational and 75 experimental spectra of PAH molecules and ions.
- Size range from 6 to 384 C atoms
- Charged states: neutral, anion (−), and cations (+ , ++, and + + +).

silicates



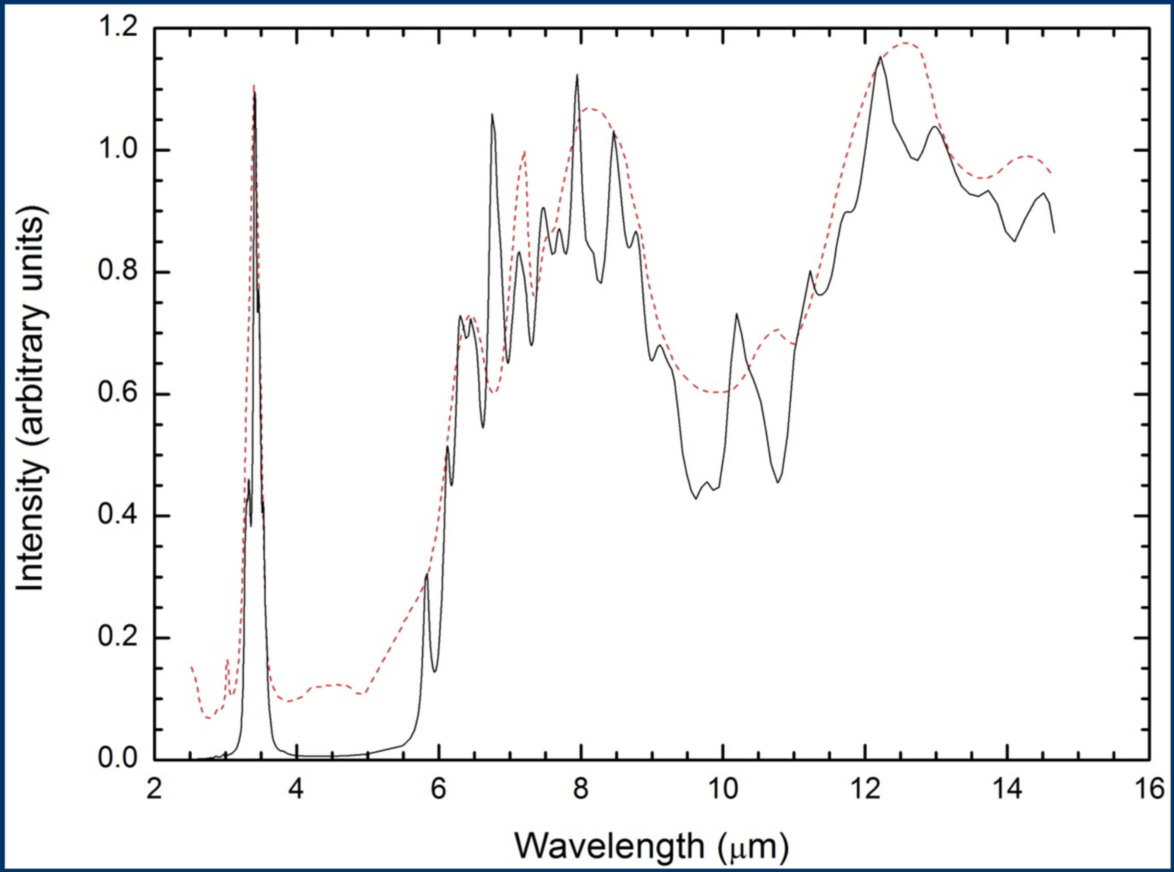
Intensity (arbitrary units)
0
5
10
15
20
25
30
35
40
6
8
10
12
14
16
18
20
Wavelength (μm)

Coal

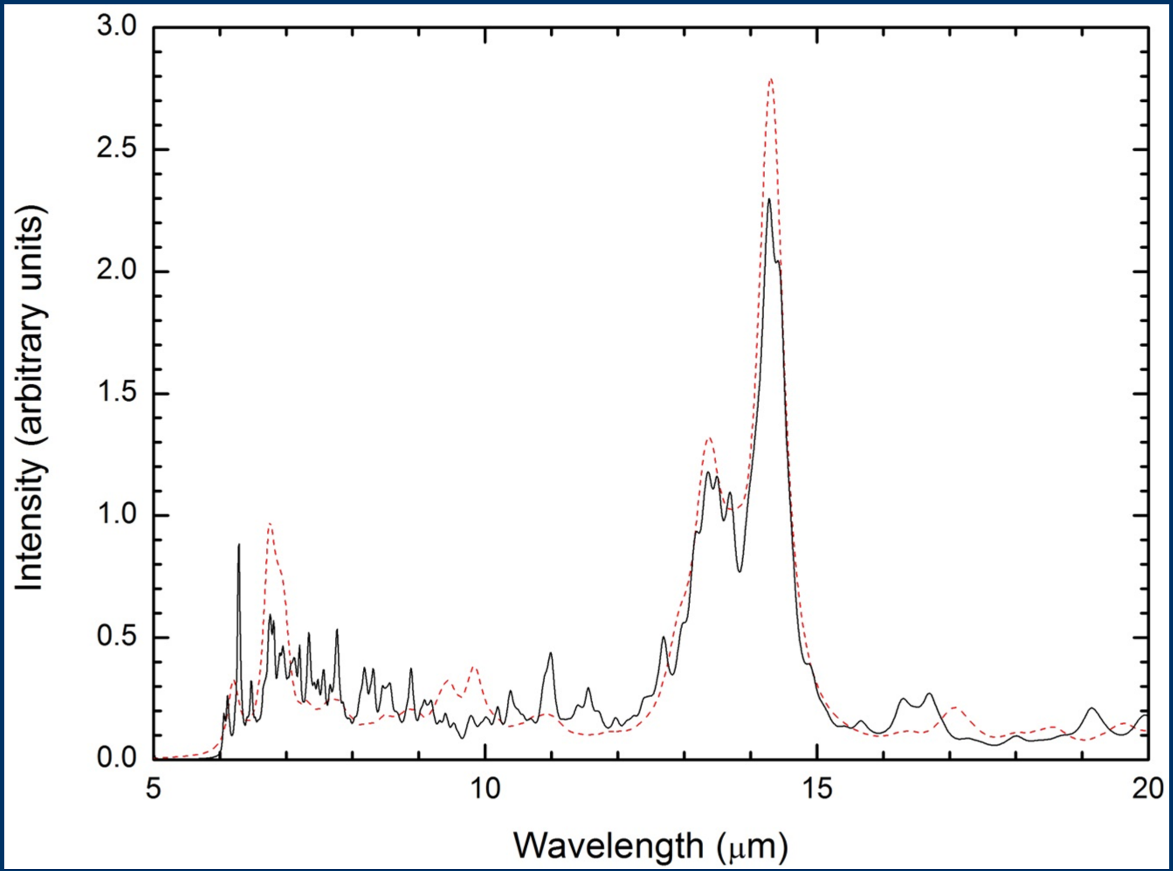


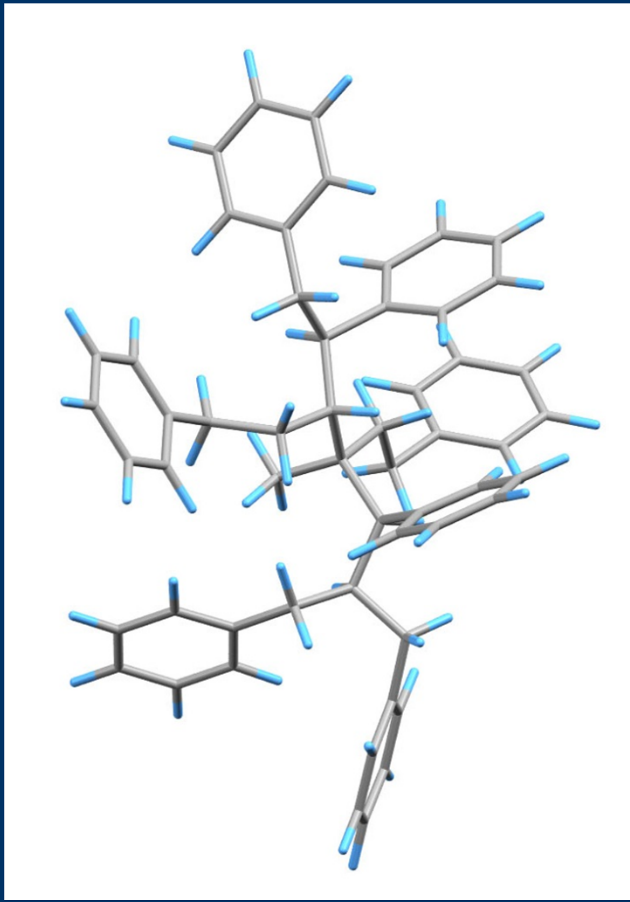
Absorbance (arbitrary units)
1.0
0.8
0.6
0.4
0.2
0.0
6
8
10
12
14
Wavelength (μm)

Hydrogenated Amorphous Carbon



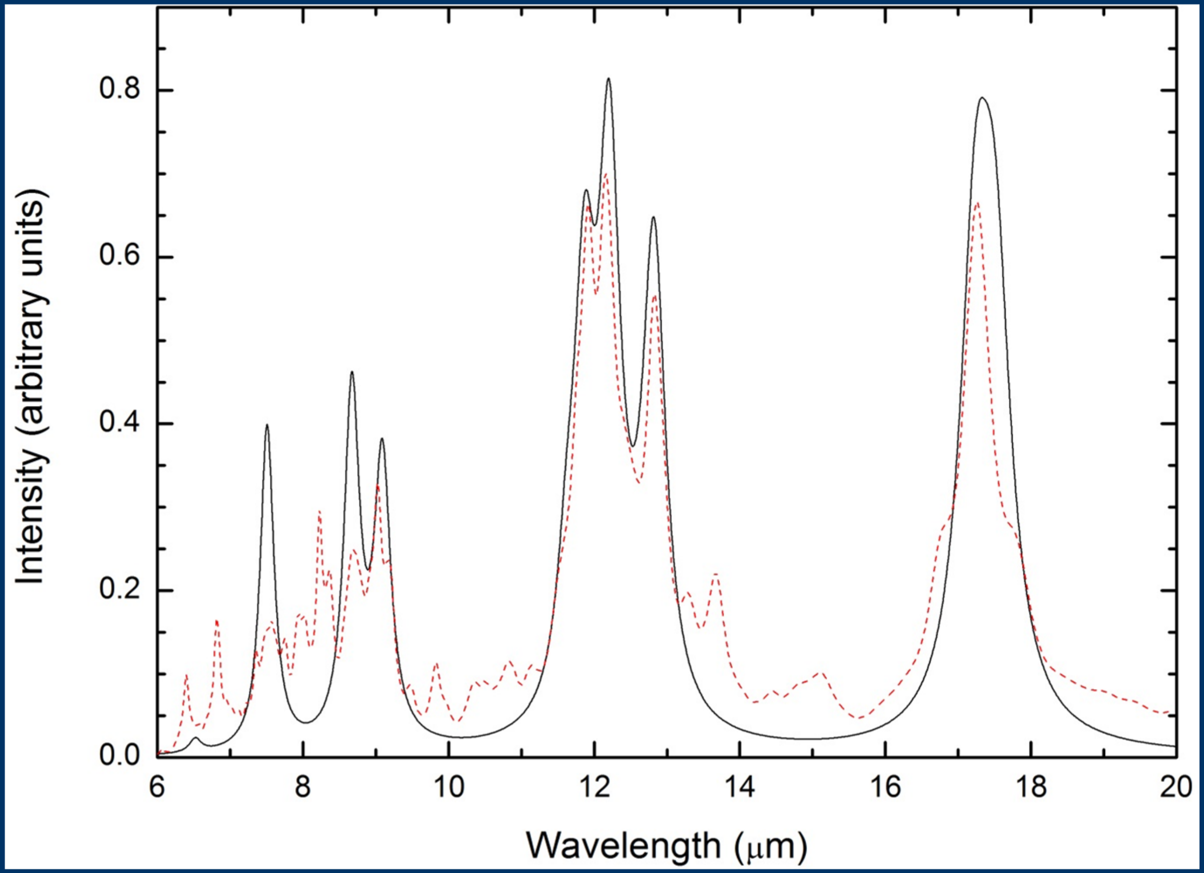
$C_{55}H_{56}$



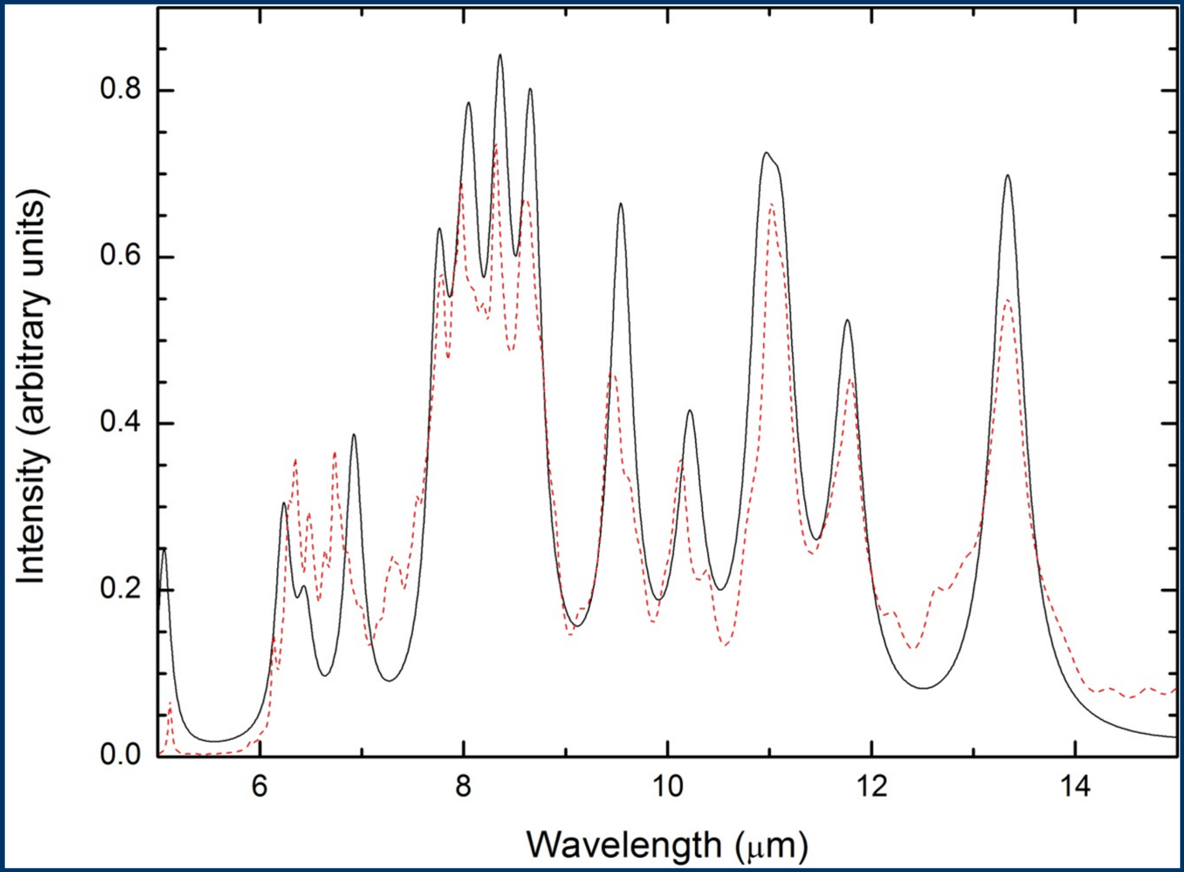


Zhang and Kwok 2015, ApJ, 798, 37

5 artificial features



10 Artificial features

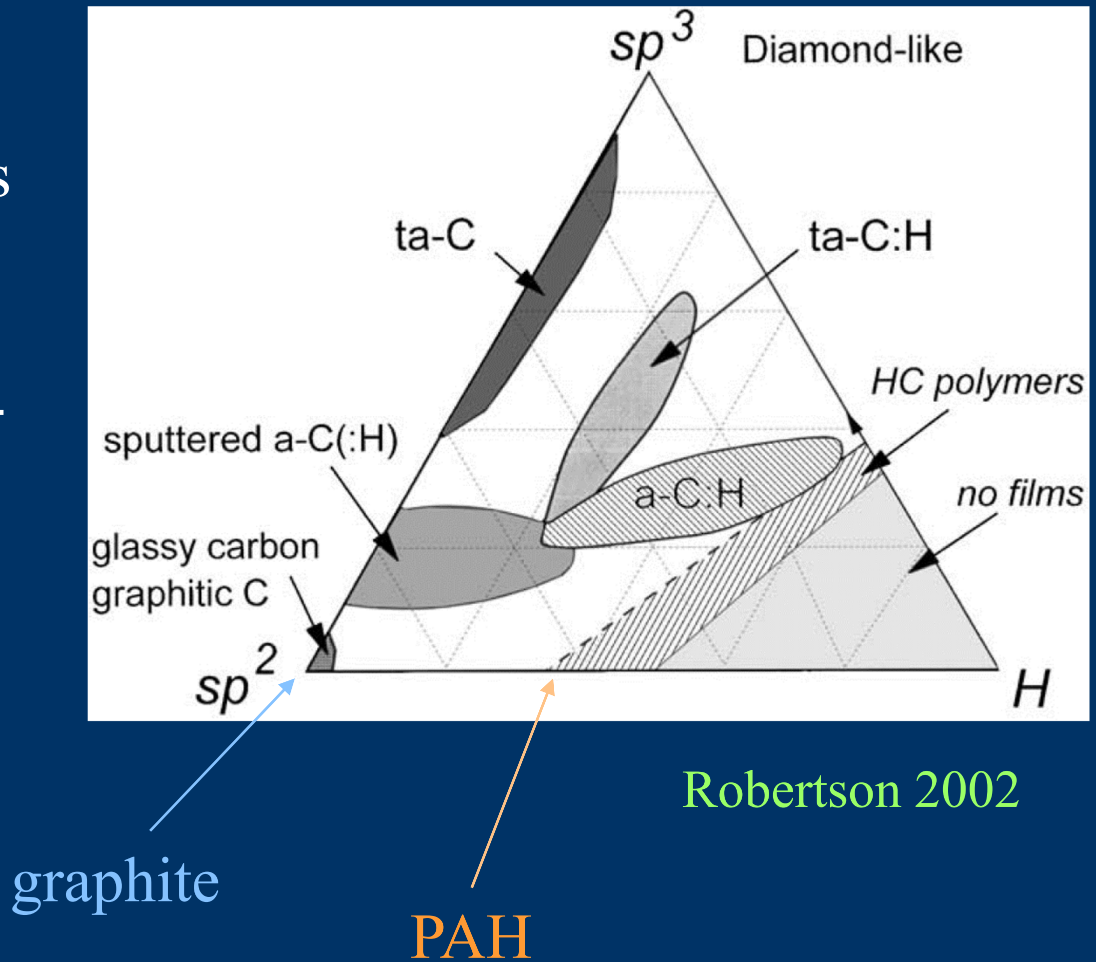


The PAH database model can fit anything!

If not PAH, then what is it?

Amorphous carbonaceous solids

- By introducing H into graphite (sp^2) and diamond (sp^3), a variety of amorphous C-H alloys can be created
- Geometric structures of different long- and short-range can be created by varying the aromatic to aliphatic ratio
- Different sp^2/sp^3 hybridization ratios, mixed hybridization states



Robertson 2002

Laboratory synthesis of carbonaceous solids

- Microwave irradiation of plasma of 4-torr methane (Sakata et al. 1987, Godard et al. 2011)
- Hydrocarbon flame or arc-discharge in a neutral of hydrogenated atmosphere (Colangeli et al. 1995, Mennella et al. 2003)
- laser ablation of graphite in a hydrogen atmosphere (Scott and Duley 1996, Mennella et al. 1999, Jäger et al. 2008)
- Infrared laser pyrolysis of gas phase molecules (C_2H_4, C_4H_6)$\Rightarrow$C-based nanoparticles (Herlin et al. 1998)
- Photolysis of methane at low temperatures (Dartois et al. 2004)
- Flame combustion forming soot (Pino et al. 2008, Carpentier et al. 2012) (C_2H_2, C_2H_4, C_3H_6 mixed with O_2)

The simplest example: Hydrogenated Amorphous Carbon

- Aromatic rings of various sizes bonded peripherally to polymeric of hydrocarbon species
- A mixture of *sp*^{2} and *sp*^{3} bonded carbon
- Formed when H content exceeds 0.1 relative to C
- Similar to soot formed from the combustion of hydrocarbons

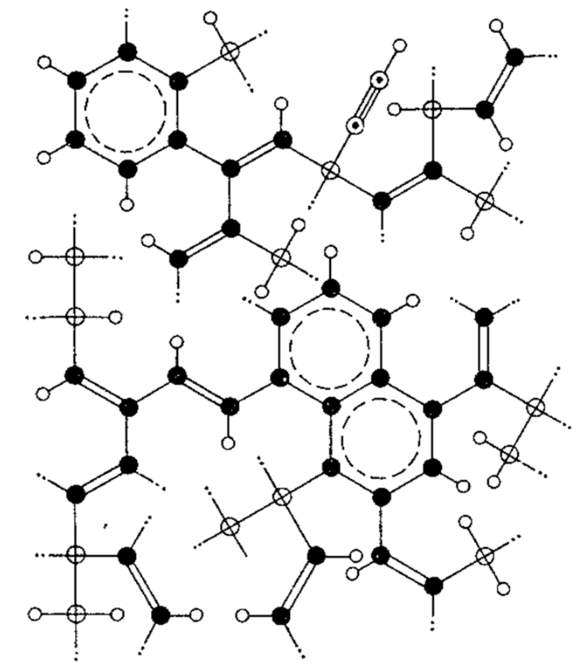


FIG. 1. Schematic diagram illustrating types of carbon atom bonding in HAC; sp^1 (circled dots), sp^2 (filled circles), sp^3 (circled crosses) and hydrogen atoms (open circles). Double dots indicate bonding to other atoms (not shown) or rare dangling bonds.

Jones, Duley, Williams 1990

Natural products of combustion

- Soot: combustion of hydrocarbon molecules in a flame
- Islands of aromatic rings connected by aliphatic chains (Chung & Violi 2011)
- Sidegroups: methyl ($–CH_3$), methylene ($–CH_2$), carbonyl (C=O), aldehydic (–HCO), phenolic (–OH), and amino ($–NH_2$) (Pino et al. 2008)
- Similar processes at work in carbon-rich circumstellar envelopes

Pure C & H or with N?

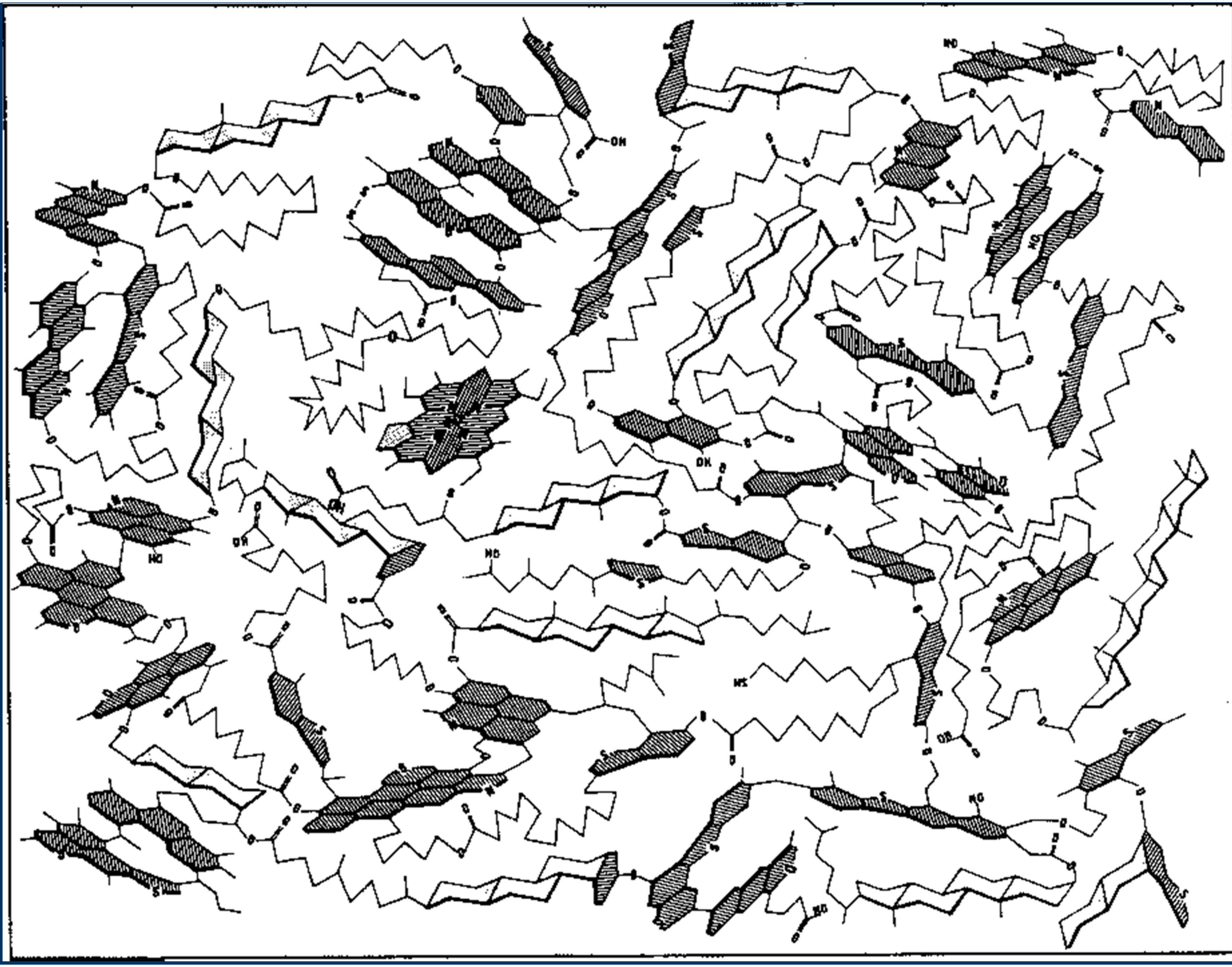
- Quenched Carbonaceous Composites (QCC): hydrocarbon plasma deposition
- Tholins: refractory organic materials formed by UV photolysis of reduced gas mixtures (N_2, NH_3, CH_4, etc.)
- HCN polymers: amorphous hydrogenated carbon nitride, formed spontaneously from HCN

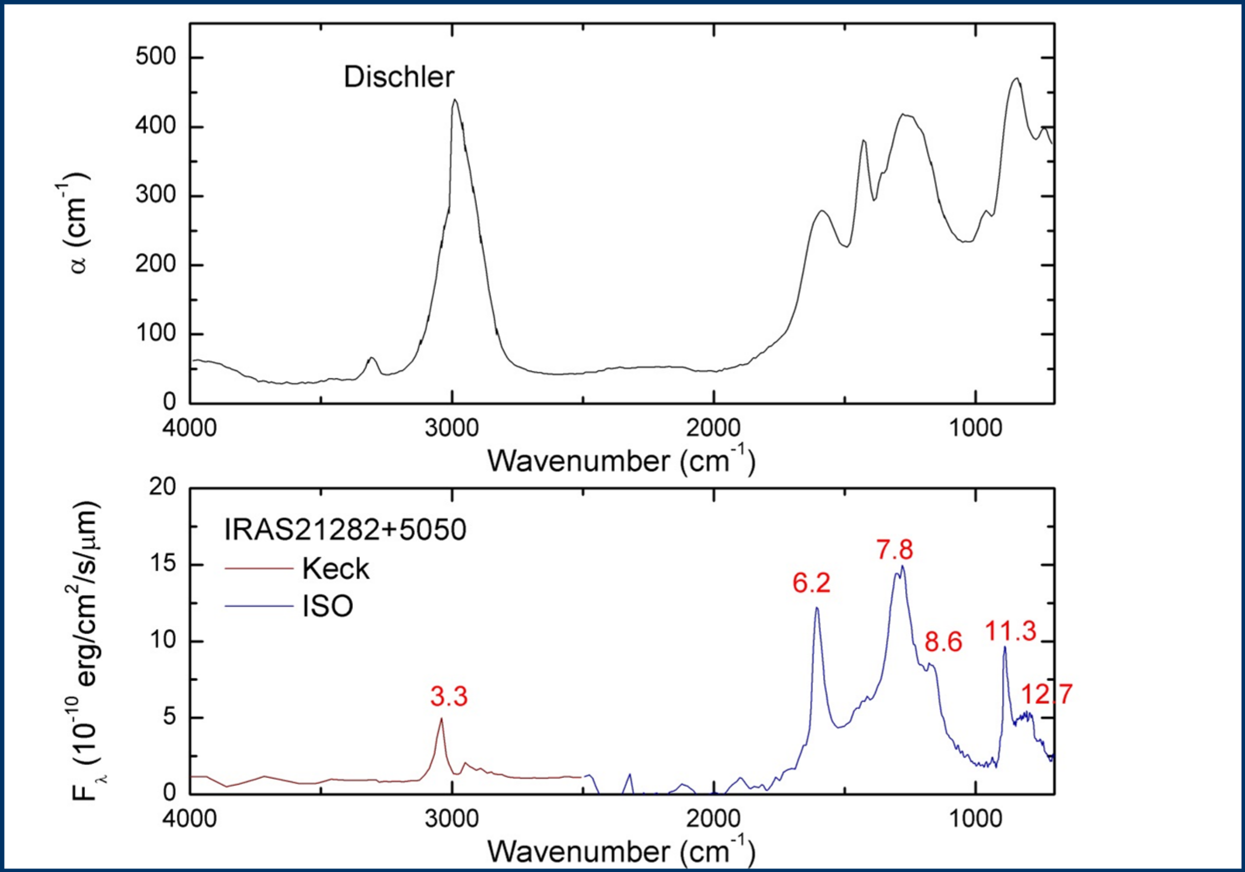
Amorphous hydrocarbon in nature

- Coal, oil, and kerogen
- Origin of kerogen: Fraction of H, S, N, and O relative to C in kerogen are similar to those in lipids $\Rightarrow$ *formed as the result of decay of living organisms*

Kerogen

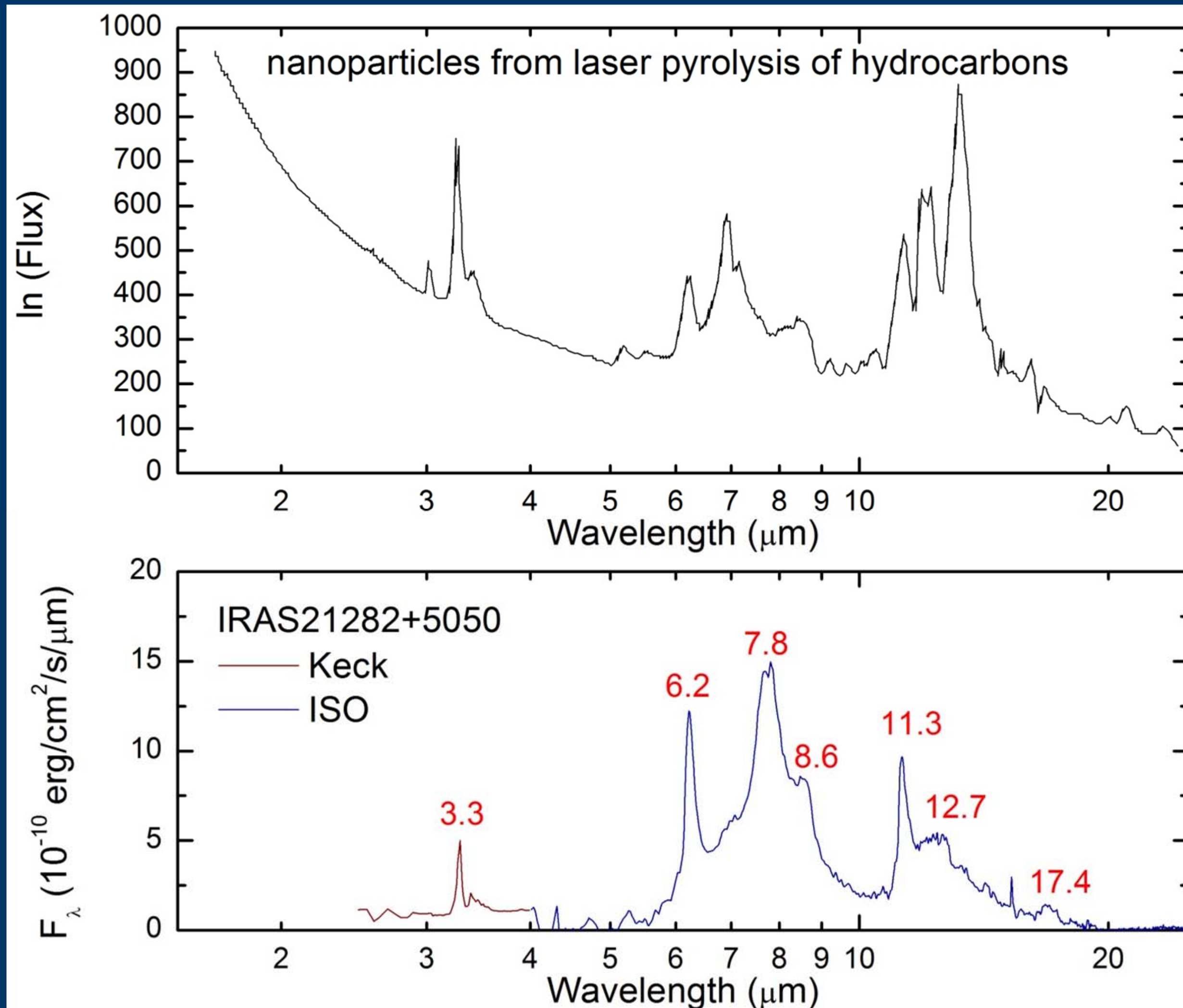
- random arrays of aromatic carbon sites, aliphatic chains $(-CH_2-)_n$), and linear chains of benzenic rings with functional groups made up of H, O, N, and S attached
- a solid sedimentary, insoluble, organic material found in the upper crust of the Earth



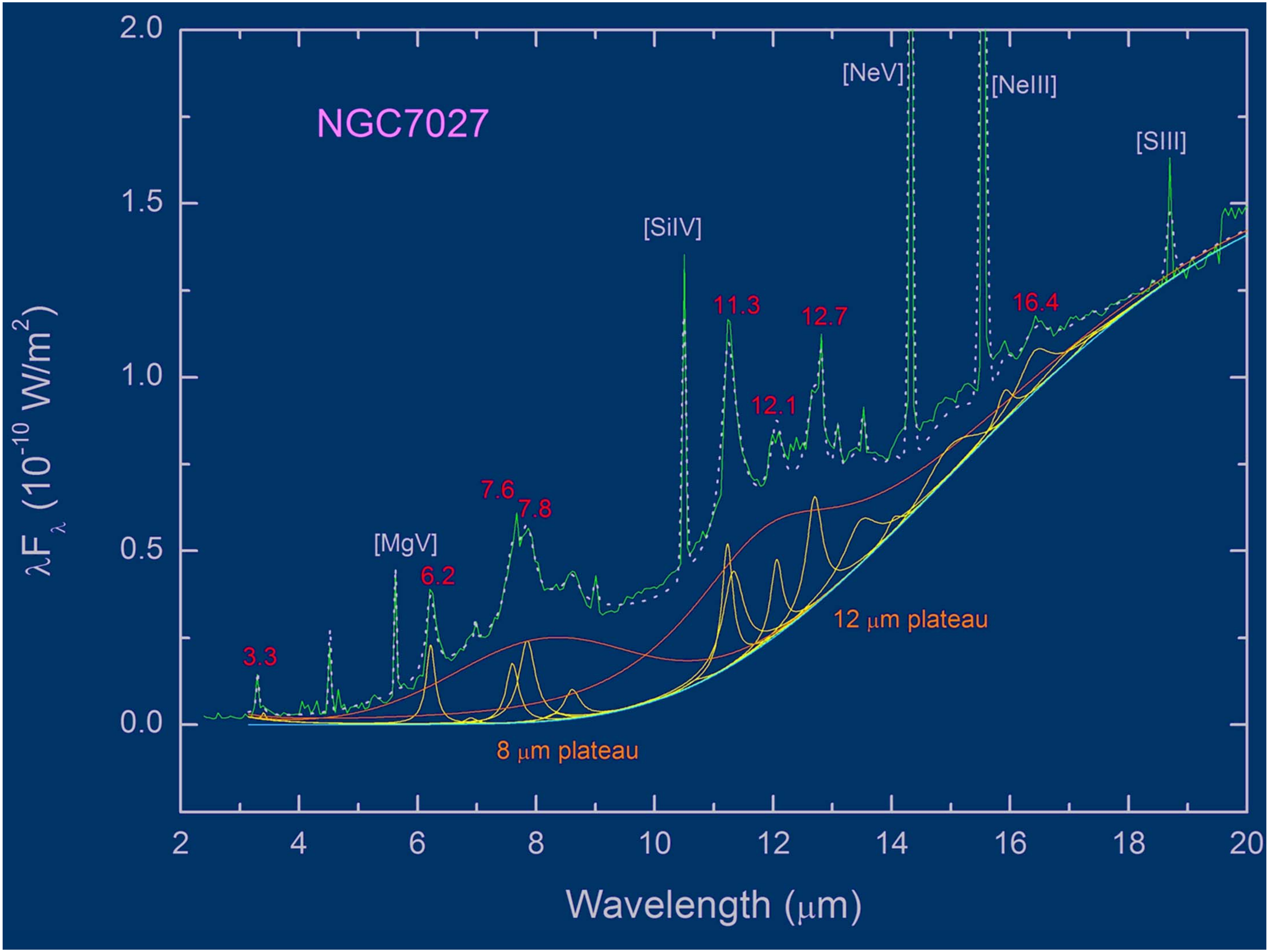


Band profiles naturally broad

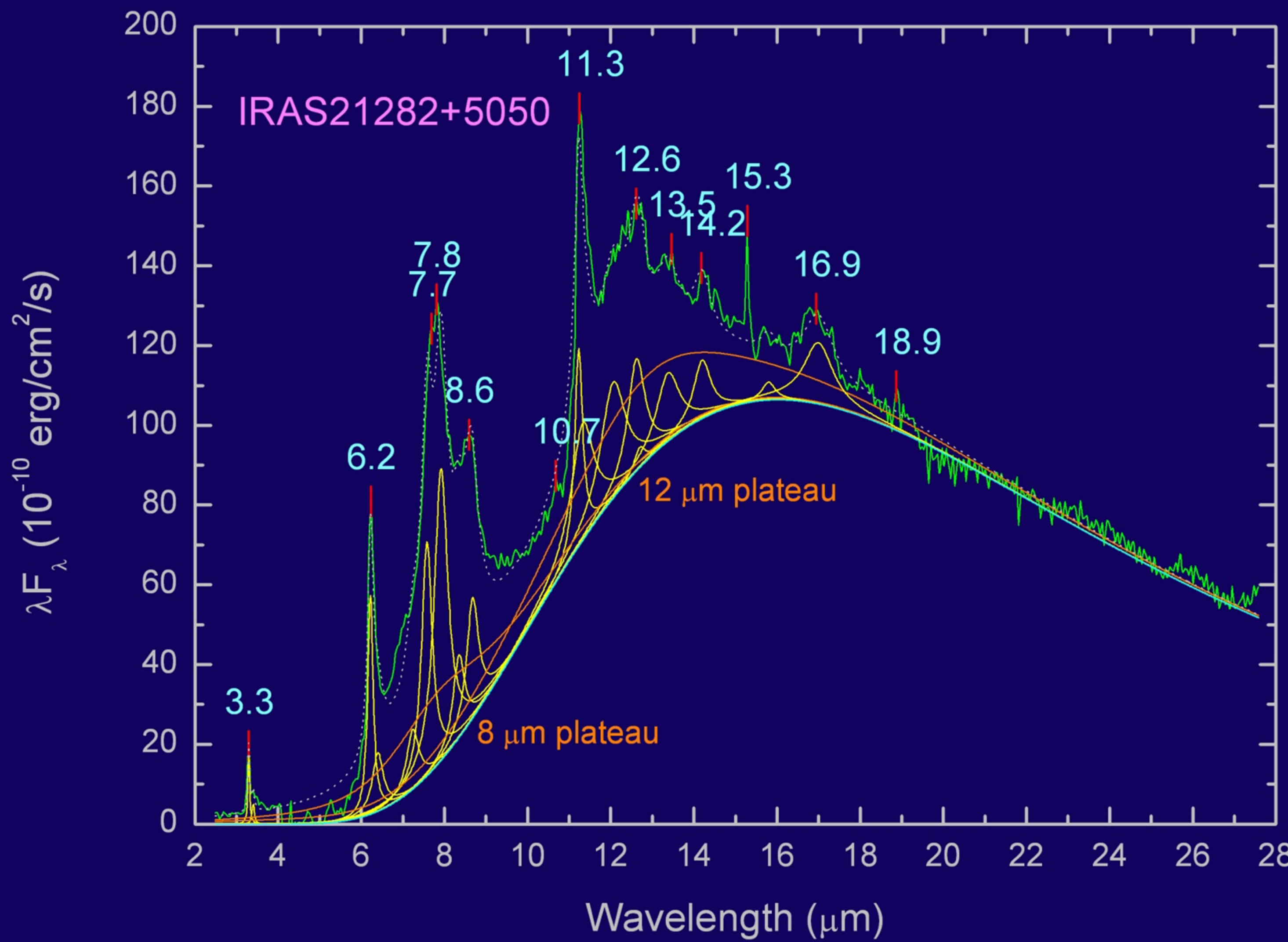
Laboratory infrared spectra of hydrogenated amorphous carbon (top, Dischler 1983) compared to the astronomical spectrum of the planetary nebula IRAS 21282+5050.

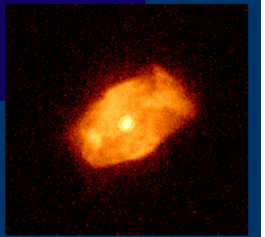


Comparison of the laboratory spectrum of nanoparticles produced by laser pyrolysis of hydrocarbons (Herlin et al. 1998) (top panel) with the astronomical spectrum of the planetary nebula IRAS 21282+5050 .



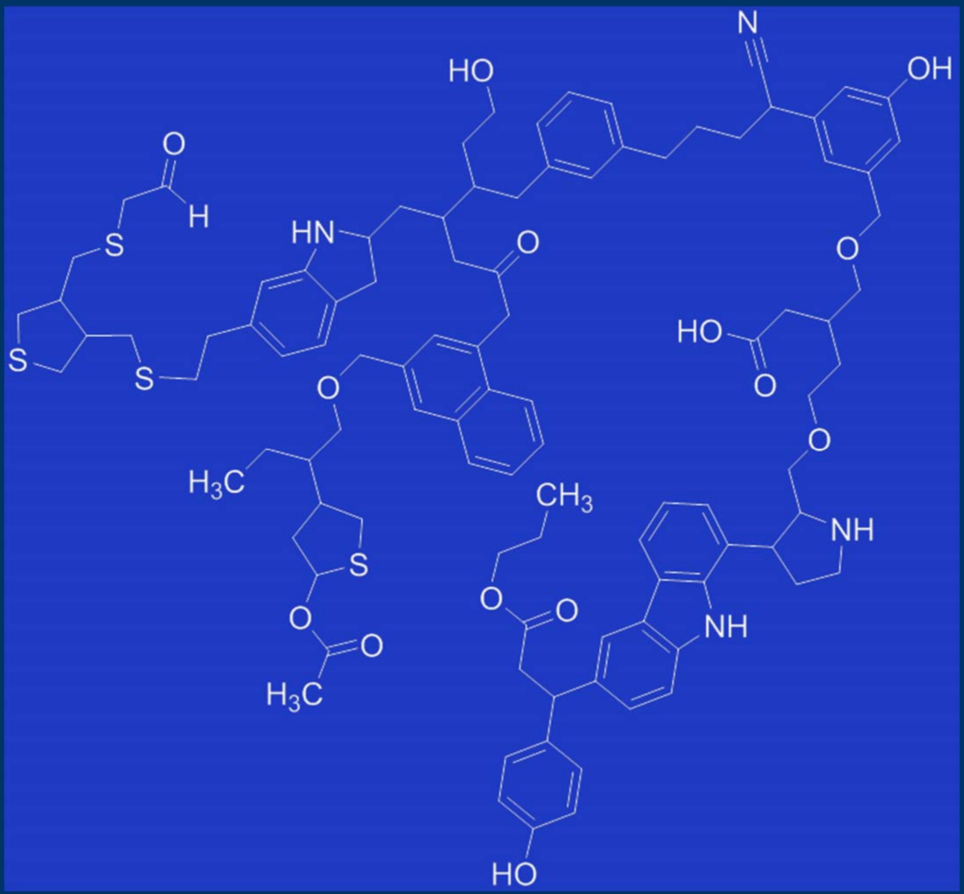
NGC7027
[NeV]
[NeIII]
[SIII]
[SiIV]
11.3
12.7
16.4
12.1
7.6
7.8
[MgV]
6.2
3.3
12 μm plateau
8 μm plateau
λF_λ (10^-10 W/m^2)
Wavelength (μm)





Flux (2.4-27.6 μm): continuum 65%, AIB: 13%, aliphatic 17%

Mixed aromatic/aliphatic organic nanoparticles (MAON) as a component of interstellar dust

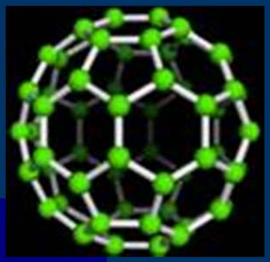


Complex organic solids with disorganized structures

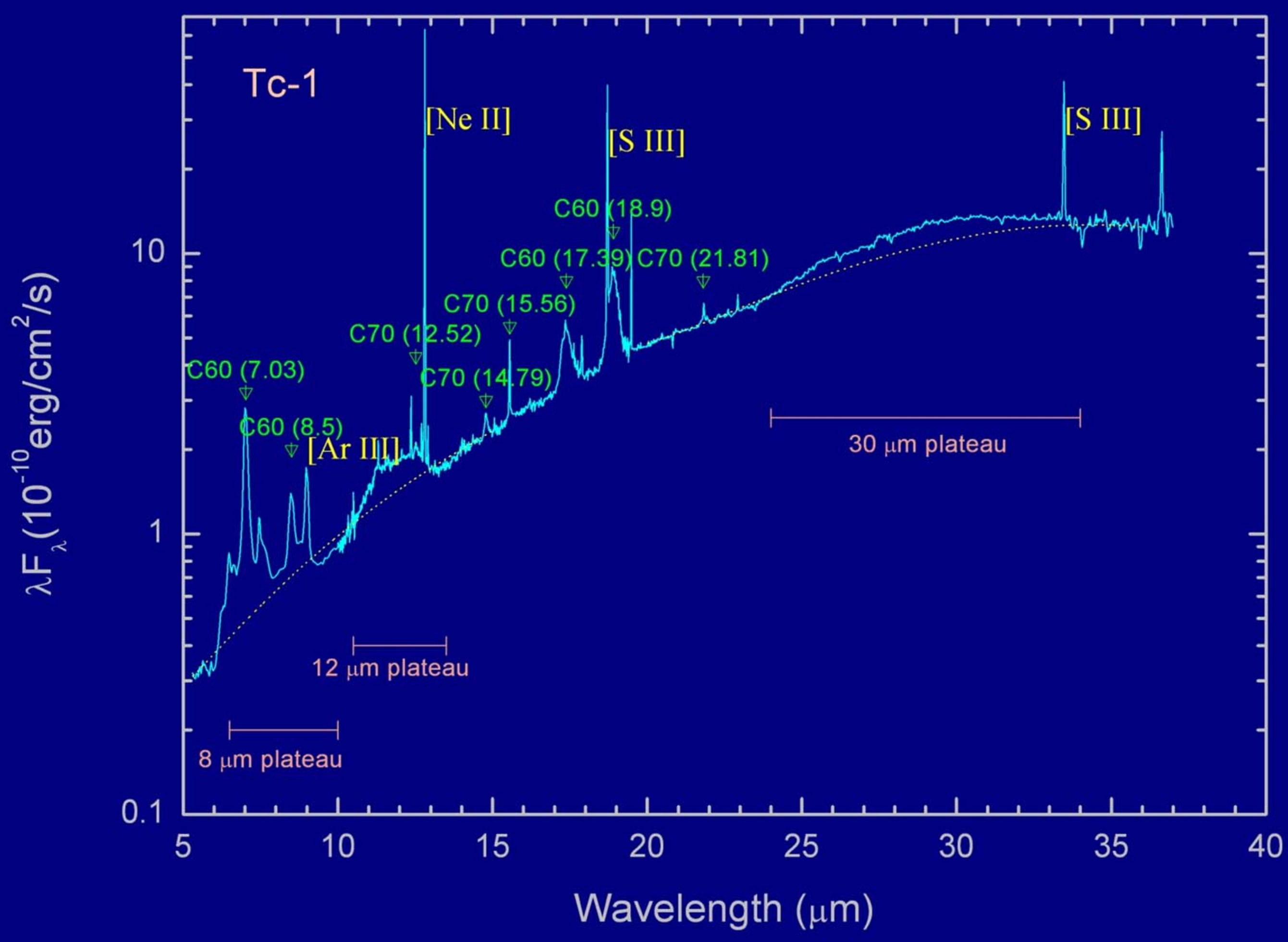
- *Small units of aromatic rings linked by aliphatic chains*
- *Impurities of O, N, S*
- *A typical nanoparticle may contain multiple of this structures*

Kwok & Zhang 2011, *Nature*, **479**. 80
Kwok & Zhang 2013, *ApJ*, **771**, 5

Fullerenes (C_{60}, C_{70})

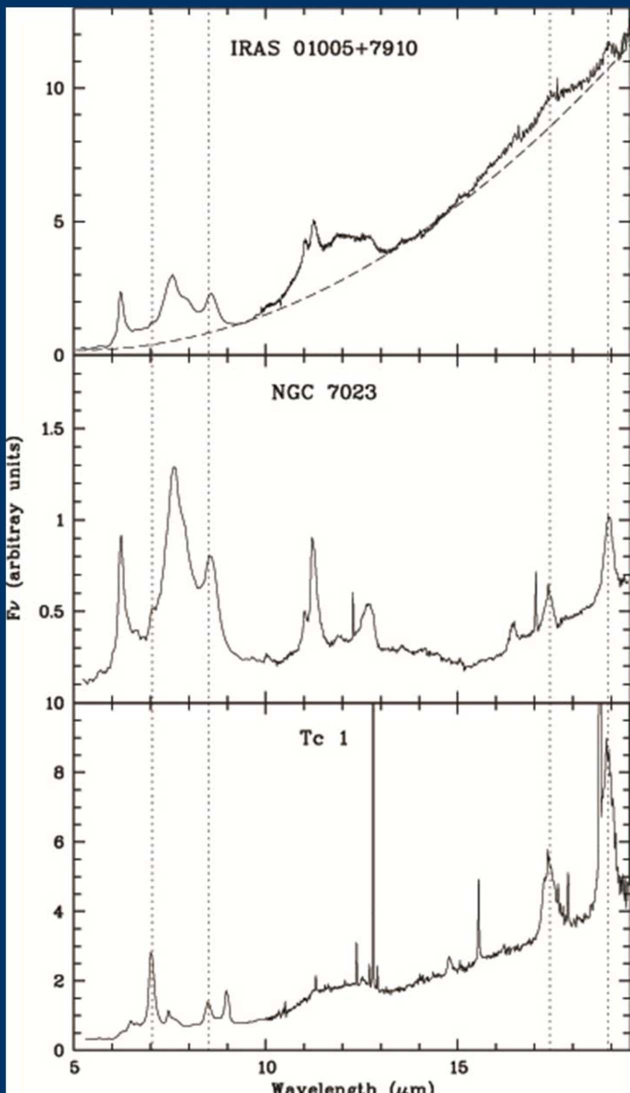


Note the plateau features



Cami et al. 2010
Sellgren et al. 2010
Garcia-Hernandez et al. 2010
Zhang & Kwok 2011

C_{60} and the plateau features



Zhang & Kwok 2011

- PPN 01005+7910: 8, 12, and 17 μm plateaus
- PN Tc-1: 8 and 12 μm plateaus
- Plateau features are present in all PN with C_{60} (Otsuka et al. 2013)
- Several C_{60}-PN (Tc-1, SMP SMC 16, SMP LMC 56) have plateau features but no AIB → not aromatic in origin
- MAONs are precursors of fullerenes? (García-Hernández et al. 2012, Bernard-Salas et al. 2012)

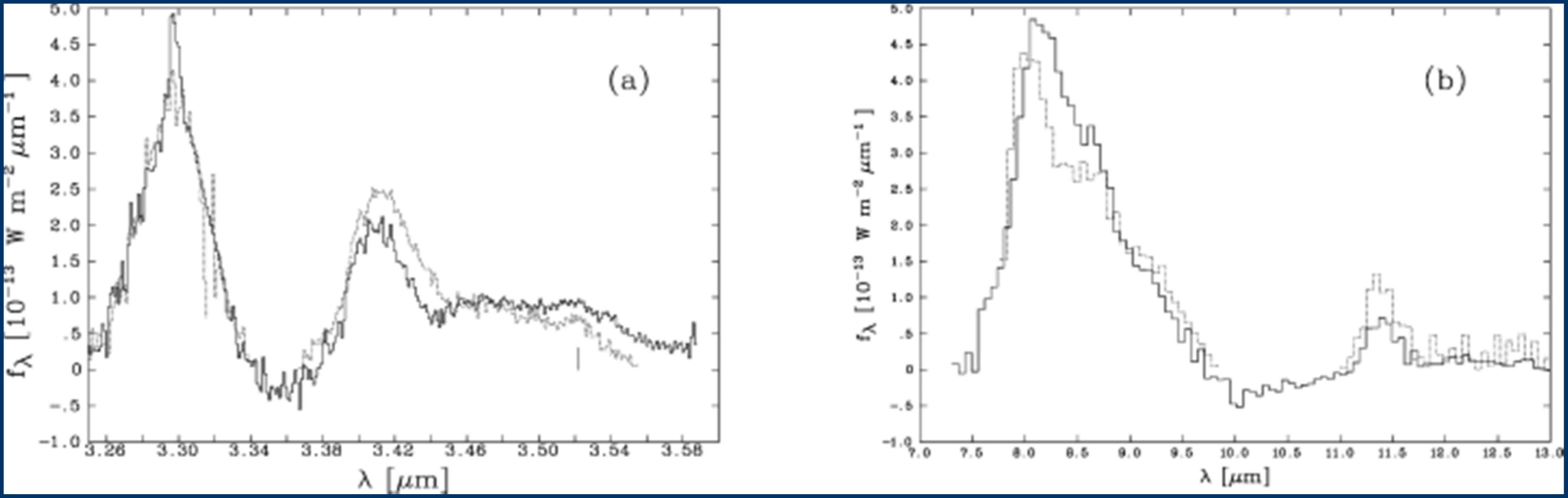
Properties of MAON

- Amorphous (no fixed structure)
- Contains rings of different sizes and chains of different lengths and random orientations
- Contains impurities
- 3-D (not 2-D)
- Exact aromatic to aliphatic ratio depends on radiation environment (photochemistry), original gas-phase components, and H content

How do they form?

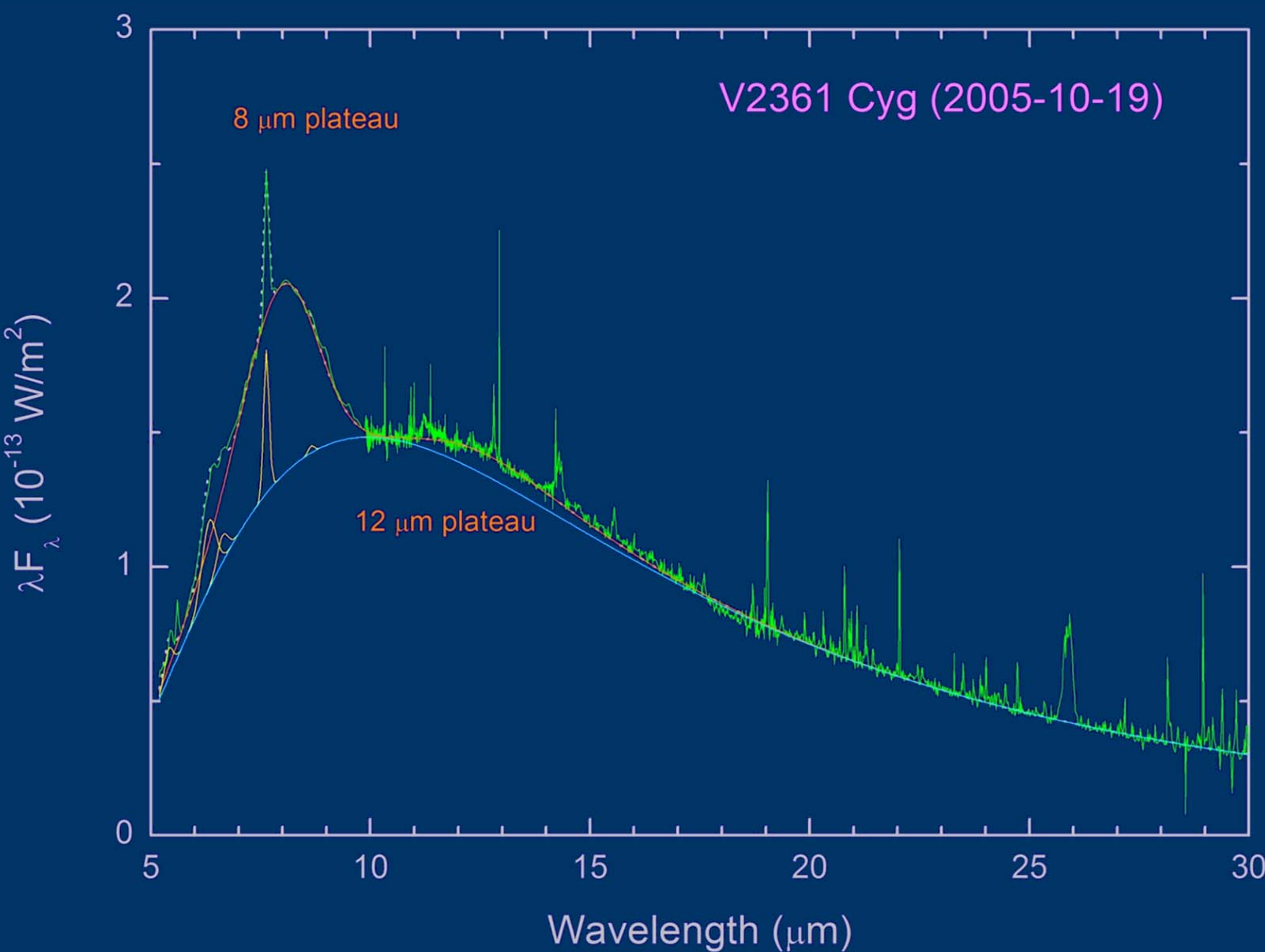
- Surface temperature of red giants: 3000 degrees
- Solid grains condensed from gas in the stellar wind under near vacuum conditions
- Theoretically impossible, especially during the PPN phase
- Observationally we see aliphatics and aromatics form in PPN on time scales as short as hundreds of years
- In novae, they form on a time scale of days

Aromatic & aliphatic features in novae



The 3.3 & 3.4 (left) and 8.2 and 11.4 μm features of Nova V705 Cas at 253 (solid line) & 320 (broken line) days after outburst (Evans et al. 2005)

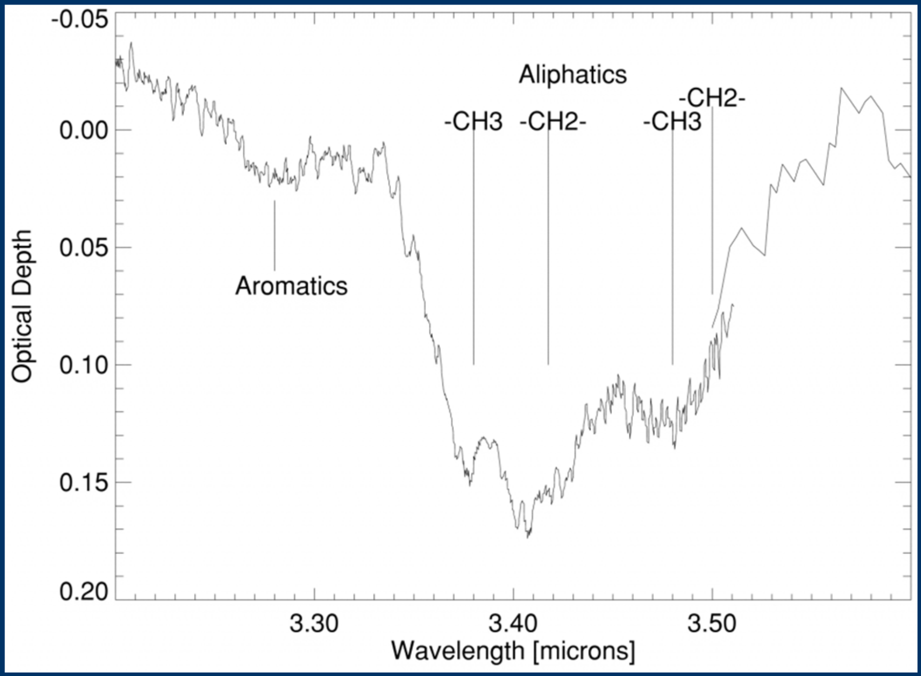
Aliphatics in novae



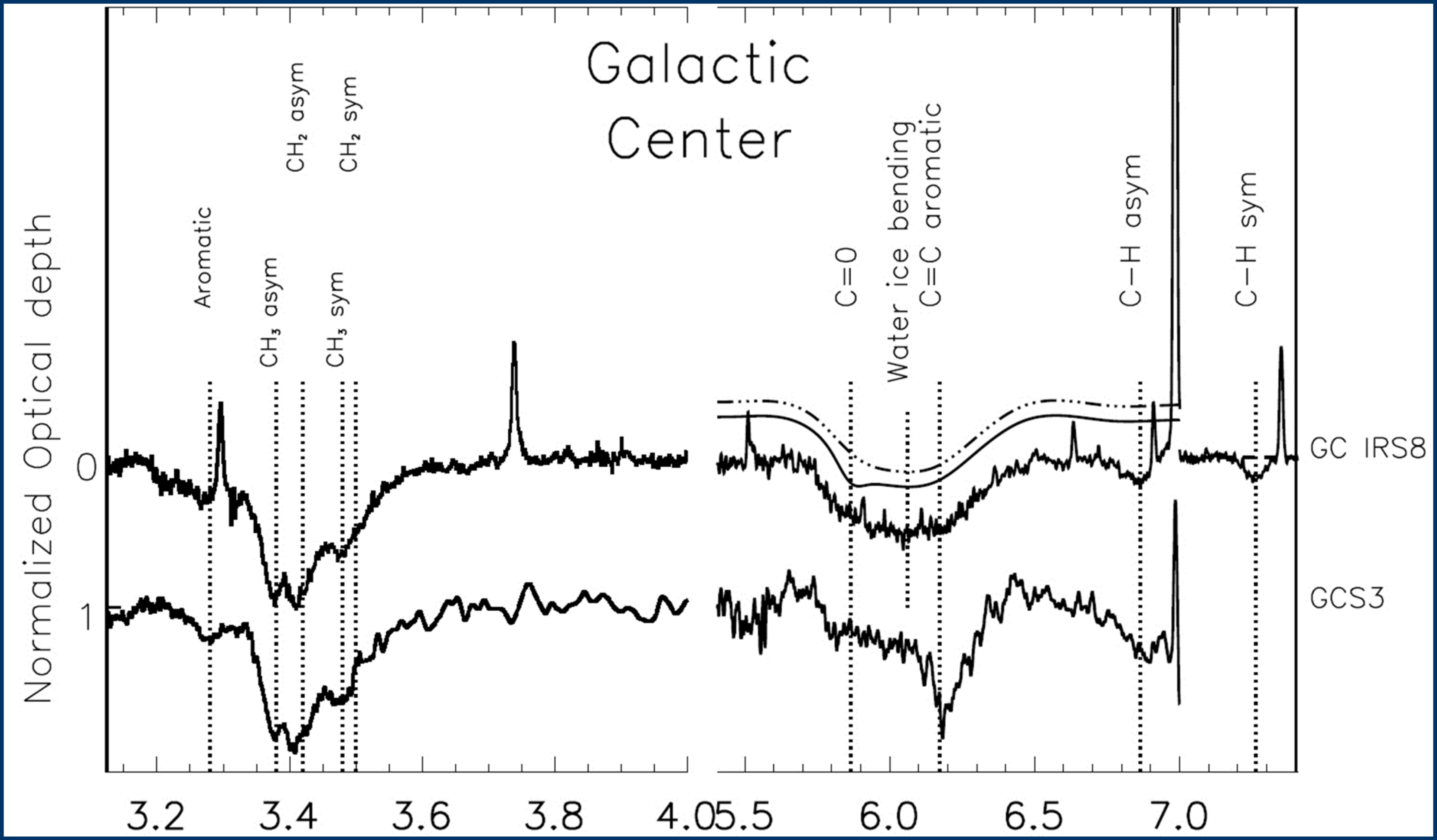
Kwok & Zhang 2011

Organic grains in the diffuse ISM

- 3.4 μm C–H stretch observed along the line of sight to the GC (Wickramasinghe & Allen 1983)
- Other sources: Sandford et al. 1991, Pendleton et al. 1994, Chiar et al. 2000
- 3.4 (stretching) and 6.9 μm (bending) aliphatic features in external galaxies: Spoon et al. 2004, Dartois et al. 2007



Aliphatics in diffuse ISM



Dartois et al. 2004,
Chiar et al. 2002,
Dartois 2011

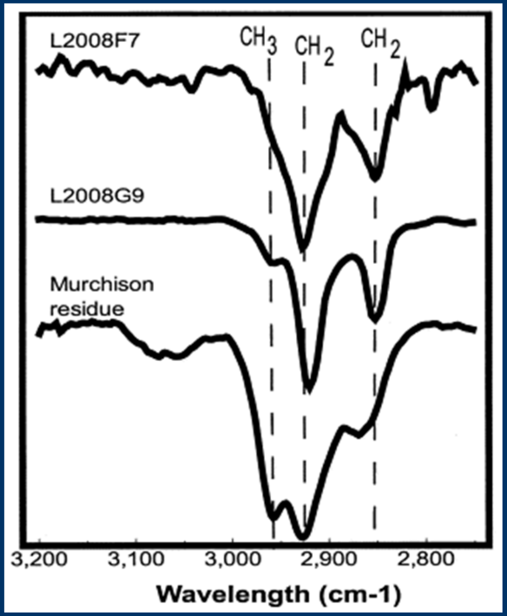
15% of C in sp^3 bonding

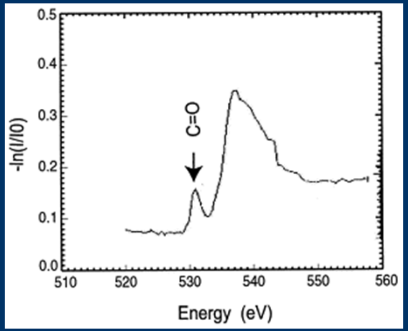
Organics in the Solar System

- Traditional picture: made up of minerals, metals, and ices
- Organic molecules and solids are found in planets and their satellites, asteroids, comets, meteorites, and minor bodies in the outer Solar System
- Origin of these organics: were they made in the solar nebula or were they inherited from interstellar space?

Interplanetary Dust

- Few microns to tens of microns in size (Brownlee 1978)
- Silicates (olivine & pyroxene)
- 10-12% carbon content
- 3.4 µm aliphatic feature and sometimes C=O group (Flynn et al. 2003)





O-XANES spectrum of IDP

The soluble component of carbonaceous chondrites

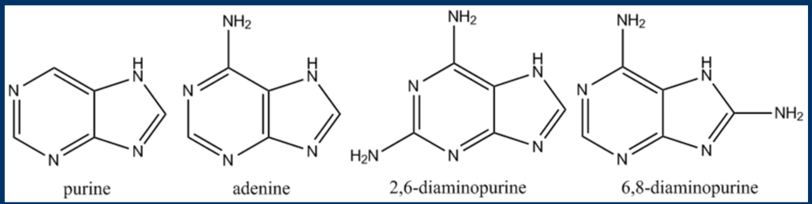
- Carboxylic acids, sulfonic and phosphonic acids, amino acids, aromatic hydrocarbons, heterocyclic compounds, aliphatic hydrocarbons, amines and amides, alcohols, aldehydes, ketones, and sugar related compounds
- >100 amino acids identified, much more than the 20 found in life on Earth
- Almost all biologically relevant organic compounds are present in carbonaceous meteorites

Schmidt-Kopplin et al. 2010, PNAS, 107, 2763

C and N isotopic ratios suggest interstellar origin (Martins et al. 2008; Nakamura-Messenger et al. 2006)

Non-terrestrial origin

- Amino acids: equal mixture of D and L chirality amino acids, non-protein amino acids not found in the biosphere, non-terrestrial values of deuterium
- Nucleobases are achiral
- Unusual nucleobases (Callahan et al. 2011, PNAS, 108, 13995).



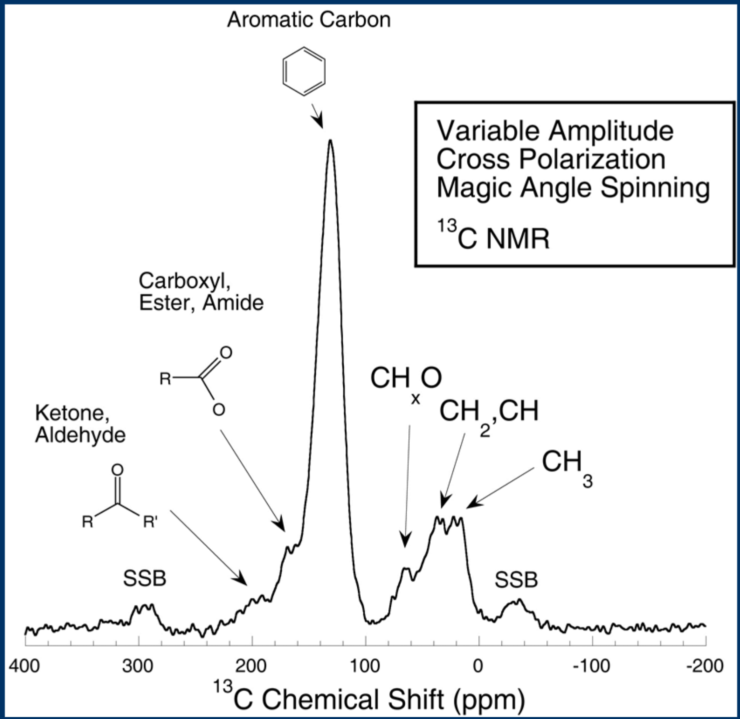
IOM in carbonaceous chondrites

- 70% of organic matter in IOM
- Destructive: thermal and chemical degradations followed by GC/MS
- Nondestrutive: NRM, FTIR, XNES, EPR, HRTEM
- Small (1-4) aromatic rings, short aliphatic chains, heteroelements (O, S, N) (Derenne & Robert 2010)
- Average abundance $C_{100}H_{46}N_{10}O_{15}S_{4.5}$ (Pizzarello & Shock 2010)

Carbonaceous condrites are products of abiotic organic chemistry: Nature can make complex organics without life

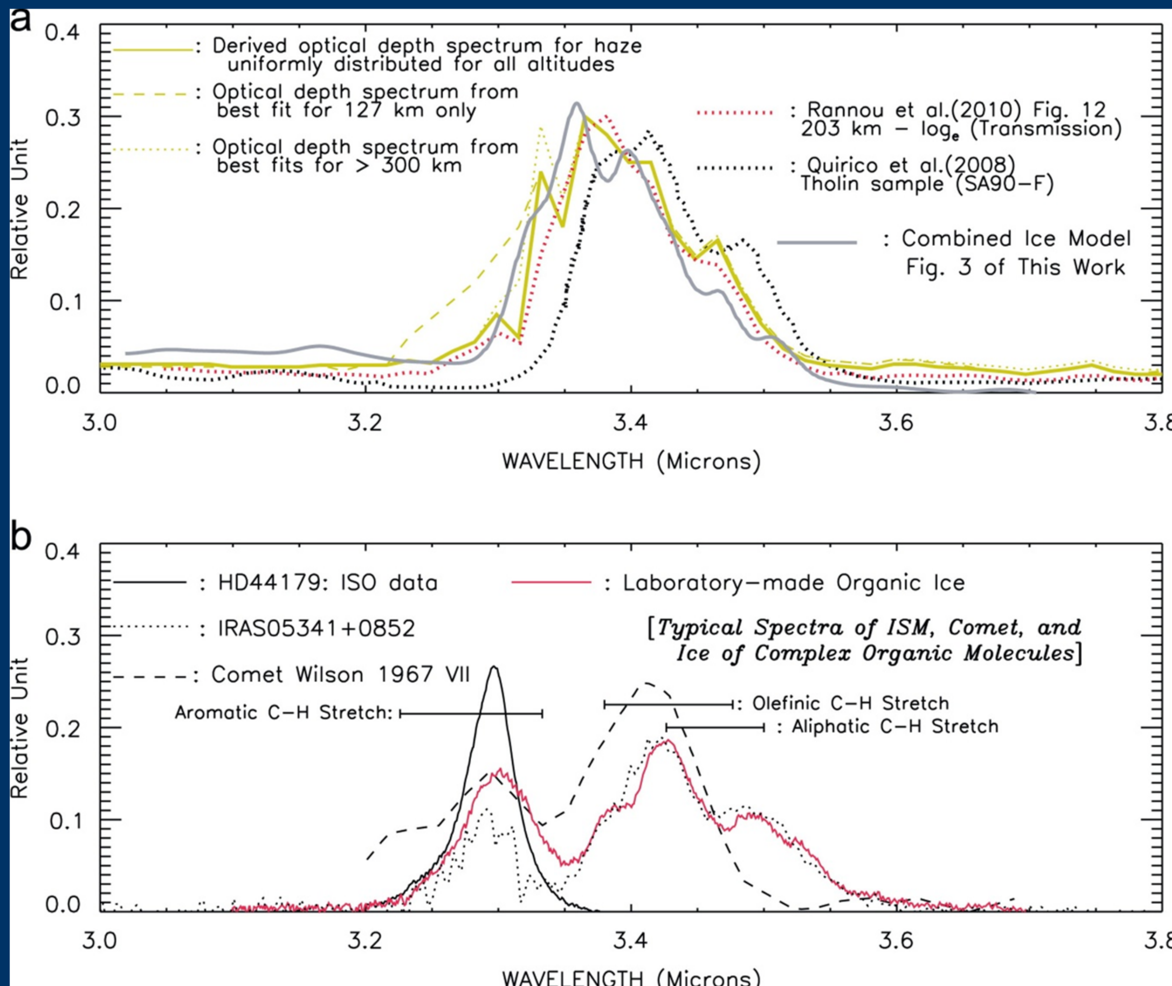
Insoluble organic matter (IOM)

- Insoluble macromolecular material similar to kerogen (Kerridge 1999)
- possibly of interstellar origin due to excess of D, ^{13}C, ^{15}N, etc.
- Composed of highly substituted single ring aromatics, substituted furan/pyran moieties, highly branched oxygenated aliphatics, and carbonyl groups (Cody et al. 2011, PNAS)



Functional groups identified in Murchison IOM (Cody 2008)

Comparison between 3.4 μm features in Titan haze, comets, and PPNs



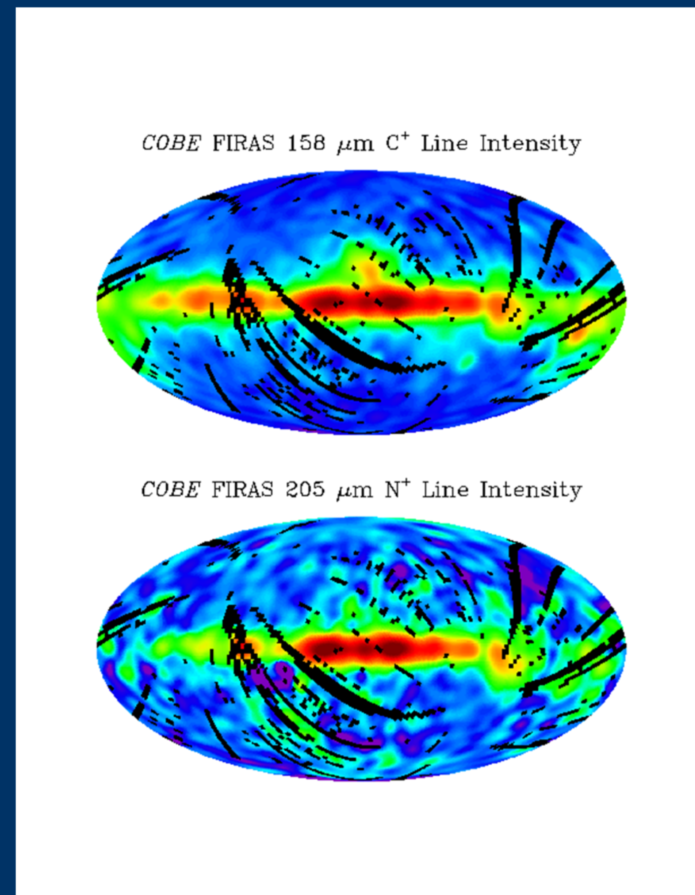
Kim et al. 2011

Future of infrared astronomy

- Distribution of neutral atoms and ions in the Galaxy
- Distribution and amount of of molecular hydrogen in galaxies
- Distribution of cold dust in galaxies
- Search for complex organics
- Distribution of organics in the Galaxy

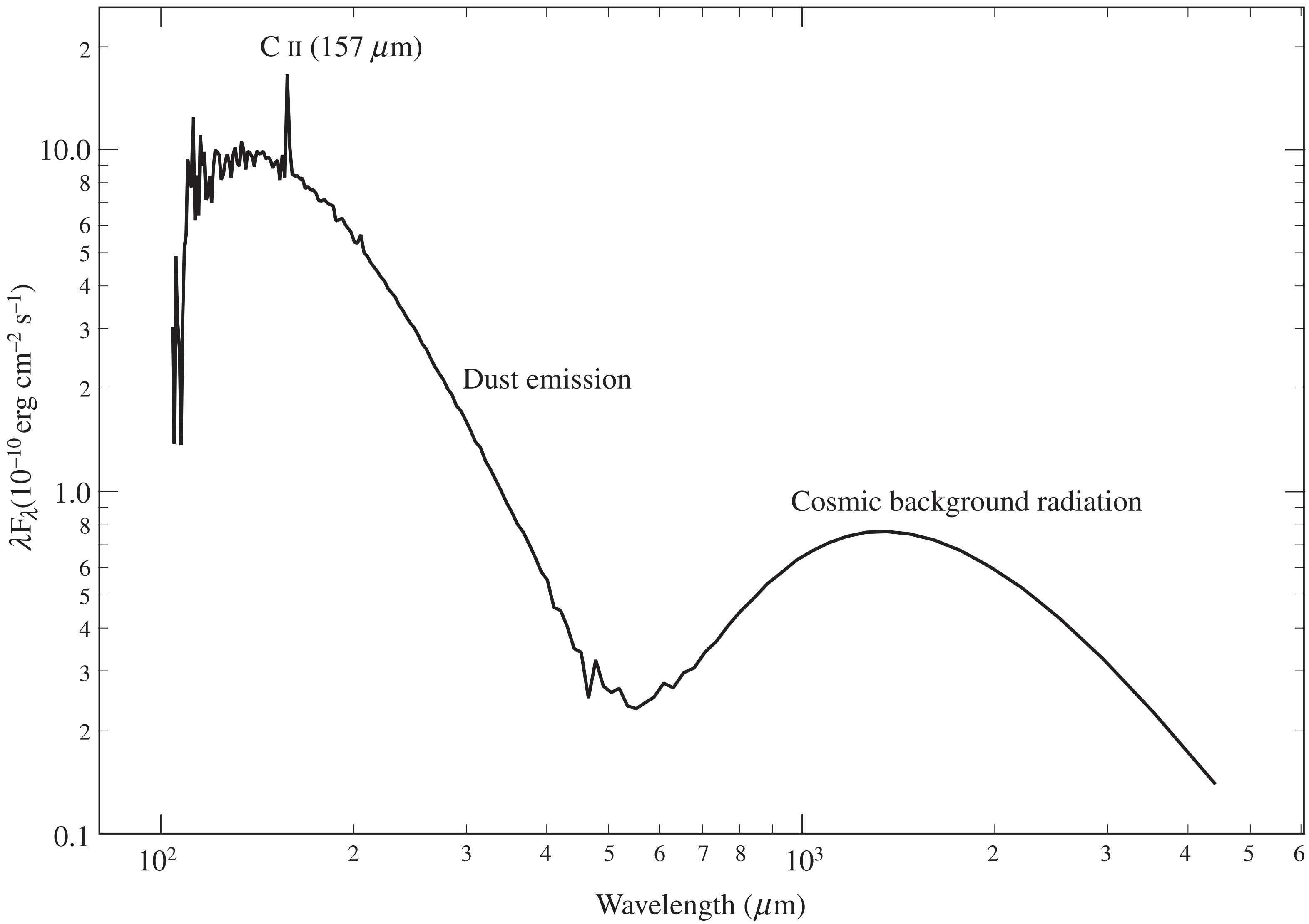
Galactic Emission

- IP of C: 11.25 eV, not protected by H from photoionization
- ${}^2P_{3/2}$-${}^2P_{1/2}$ (158 μm)
- Single brightest emission line of galaxies (total luminosity of CII line: 5×10^7 L◆)

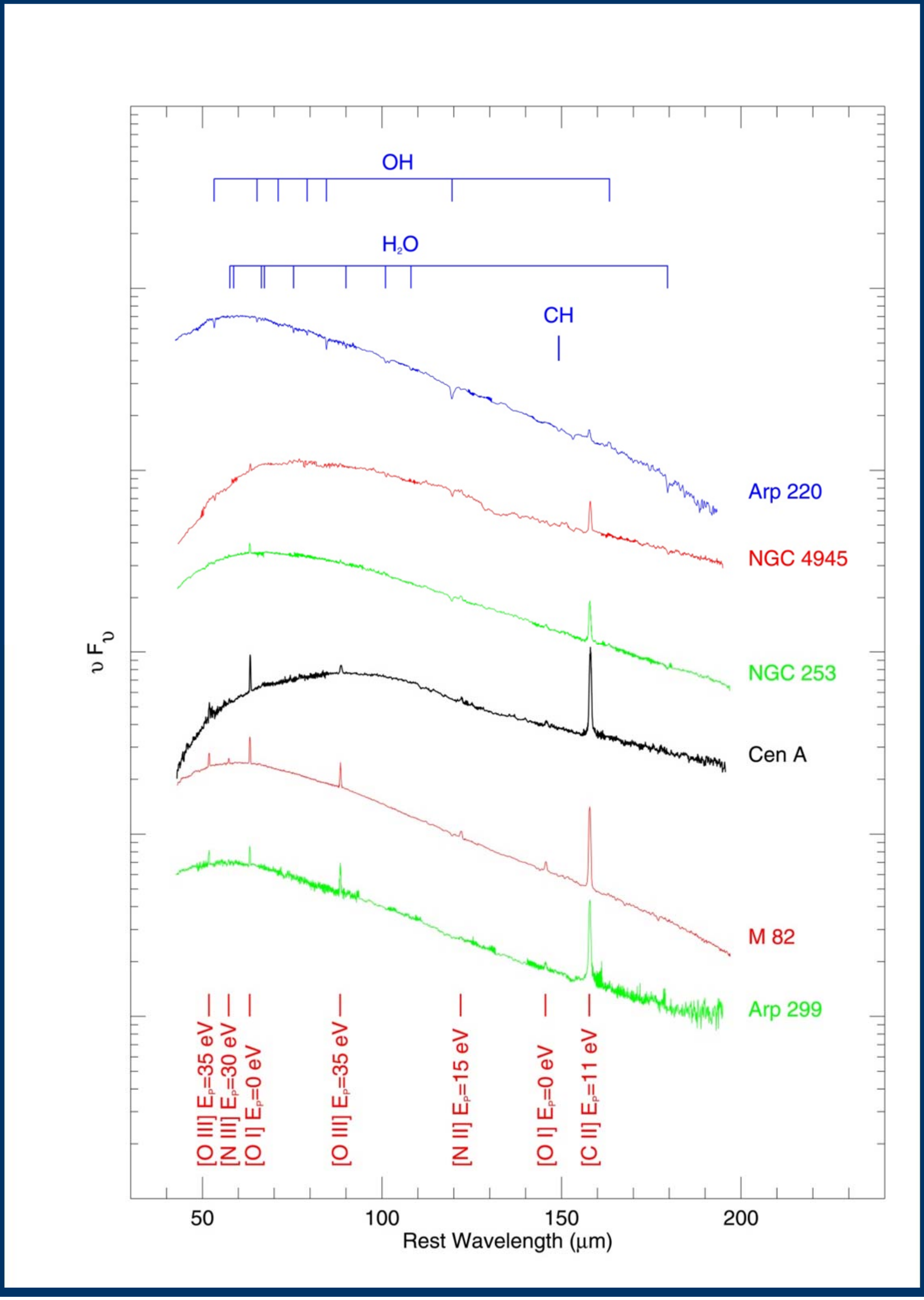


Wright et al. 1991

0.3% of total dust luminosity



Integrated spectrum of the galactic plane as observed by COBE



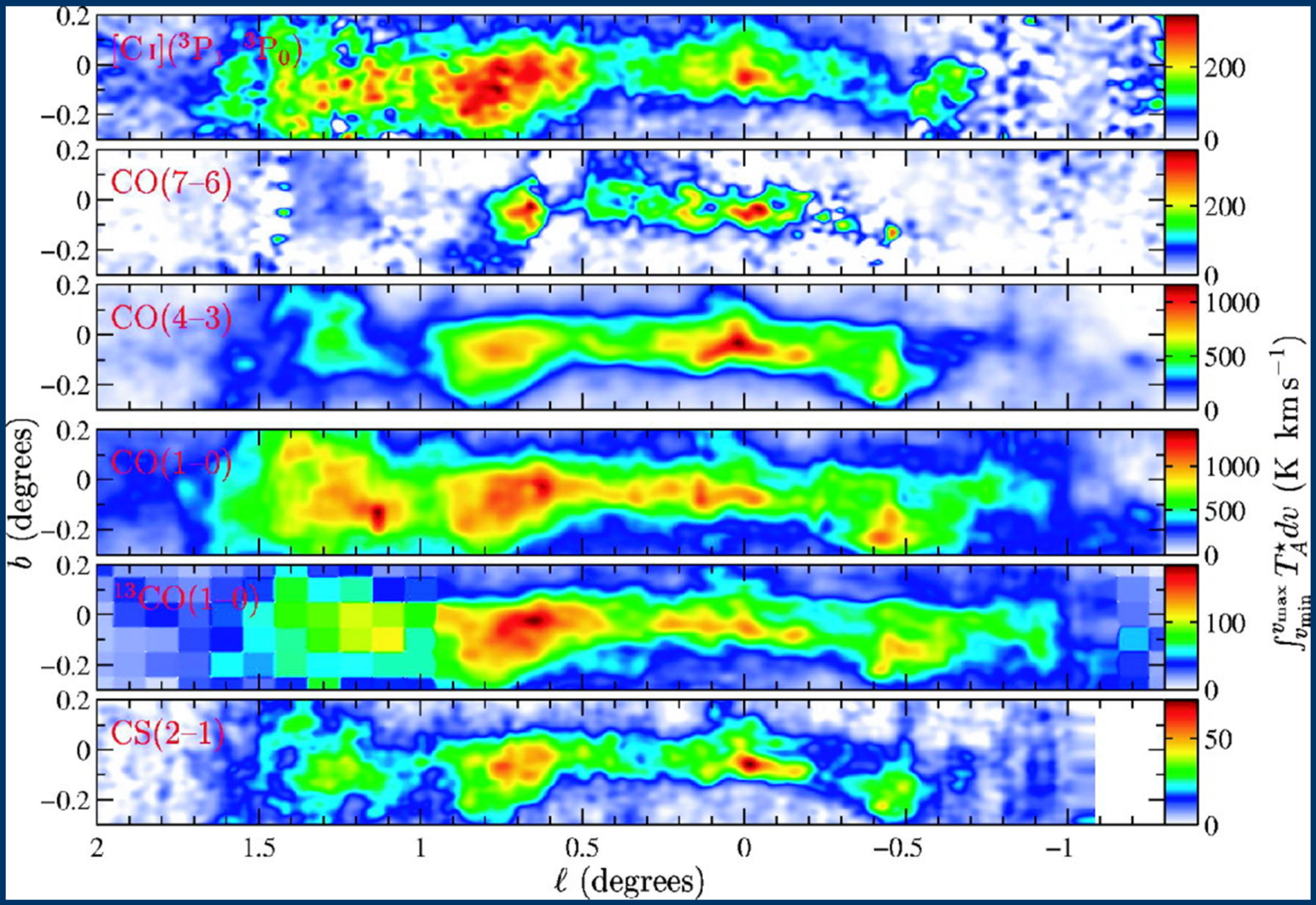
[CII] is often the single brightest emission line in the spectrum of galaxies

A wide field narrow band survey

- A 1-m telescope at 300 μm will have ~1 arcmin resolution
- Survey of the galactic plane in [CI] (370, 609 μm), [CII] (158 μm), [NII] (122, 205 μm)

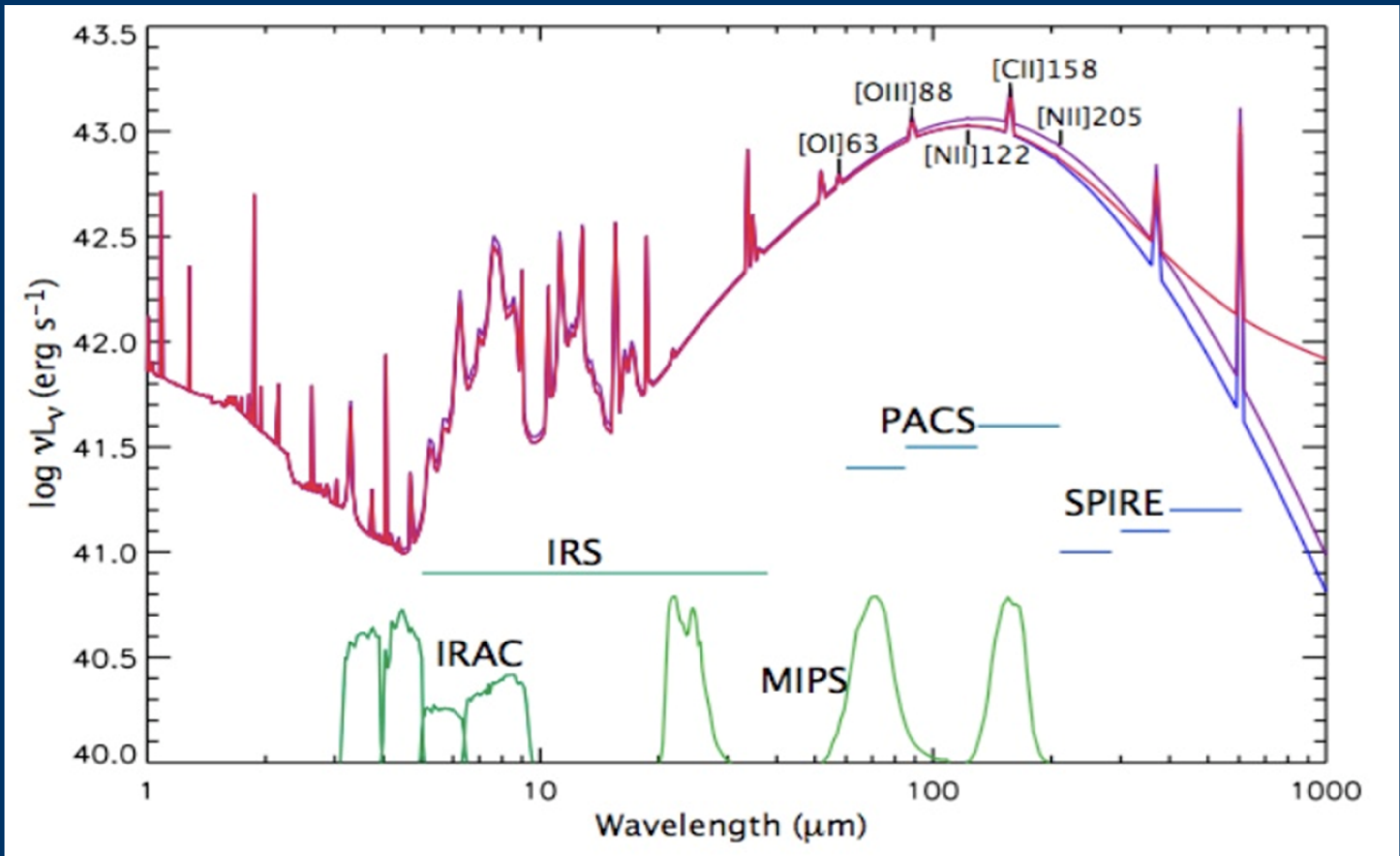
Not affected by dust extinction, can trace the structure of the Galaxy

Inner 3º of Galactic Center mapped by AST/RO 1.7 m



Martin et al. 2004

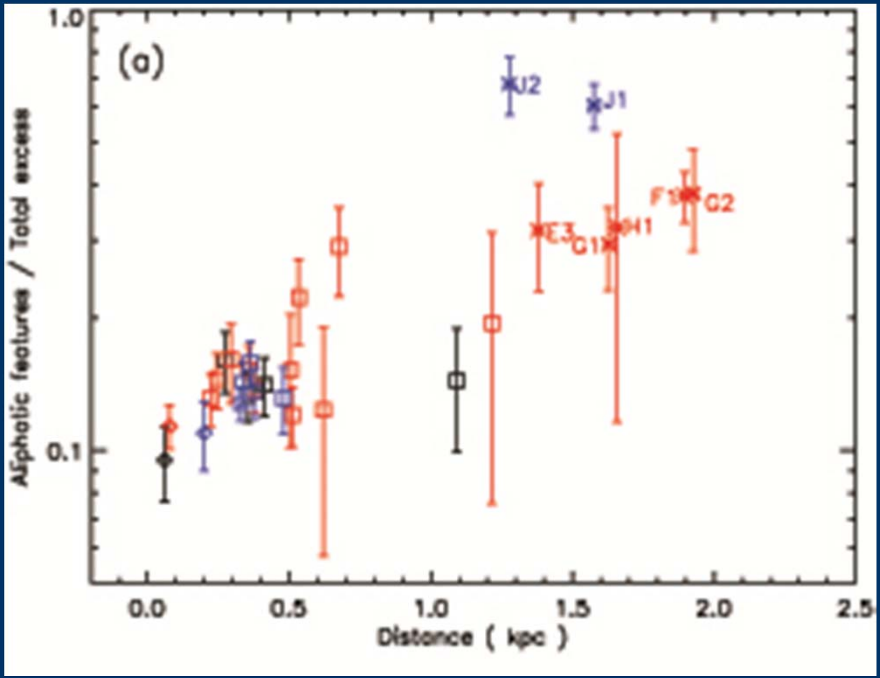
SED of galaxies

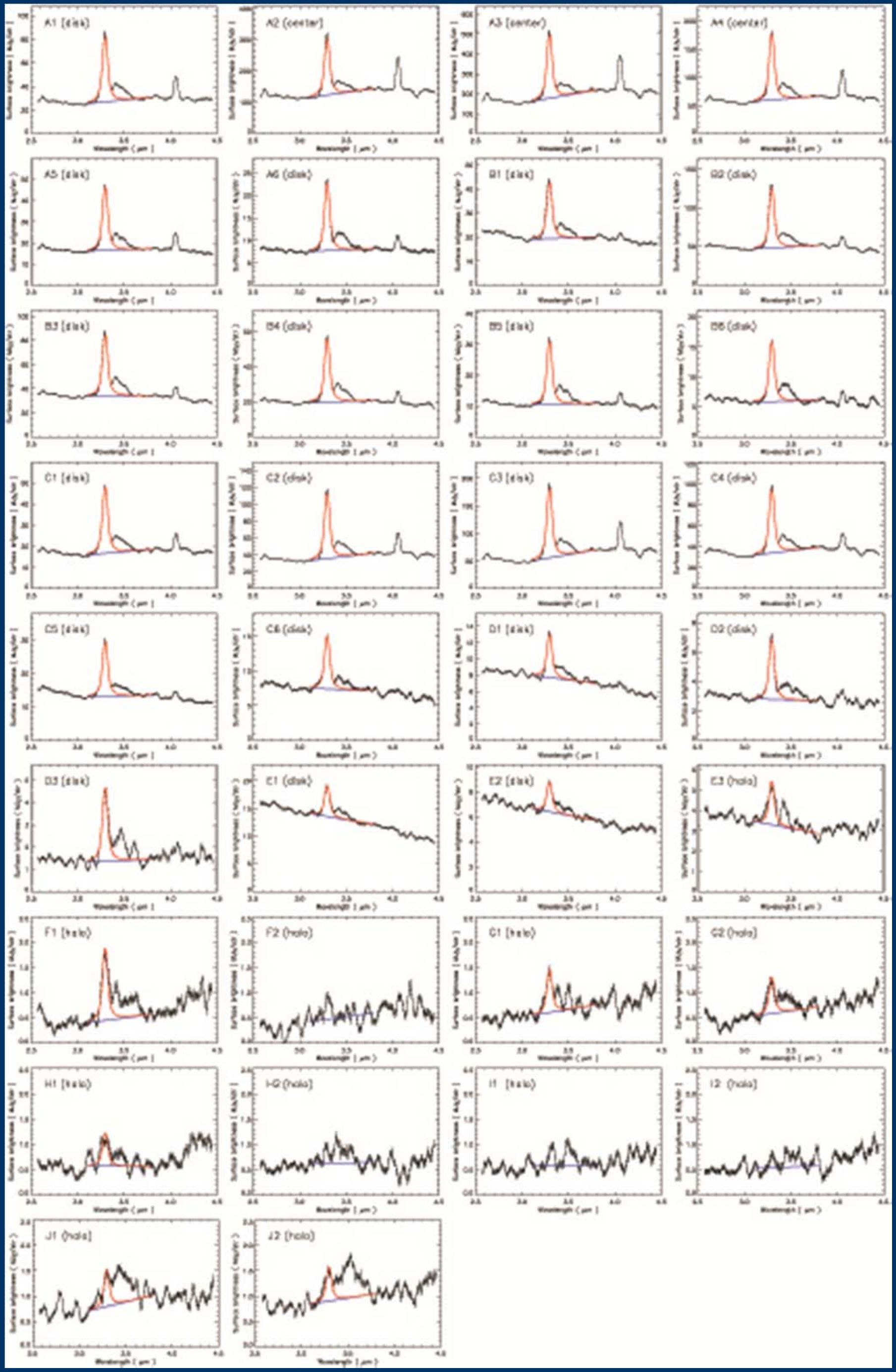


Mapping of the aliphatic component of the Galaxy

- At least 15% of C is in the form of aliphatic compounds
- They are everywhere in the diffuse ISM

- Akari observations of M82
- 3.4 µm feature strength increases from disk to halo





Yamagishi et al. 2012

What can a small telescope from Antarctica do?

Summary

- Organic compounds are everywhere in the Universe (from solar system to ISM to galaxies)
- Detection of AIBs in galaxies with $z \sim 2$ suggests that complex organics already present in the early history of the universe
- The observed features are consistent with the carrier being mixed aromatic/aliphatic organic nanoparticles (MAON)
- PPN, PN, and novae are the only objects that we have direct observations of organic synthesis
- Chemical evolution leading to complex organic compounds can take place over only a few thousand years in the circumstellar environment

Complex organics are routinely made by ordinary Sun-like stars

References

Kwok, S. 2004 The Synthesis of Organic and Inorganic Compounds in Evolved Stars, *Nature*, **430**, 985

Kwok, S. & Zhang, Y. 2011, Mixed aromatic/aliphatic organic nanoparticles as carriers of the unidentified infrared emission features, *Nature*, **479**, 80

Kwok, S. 2011, *Organic Matter in the Universe*, Wiley